GALACTIC PEVATRONS: LINKS WITH COMPACT OBJECTS

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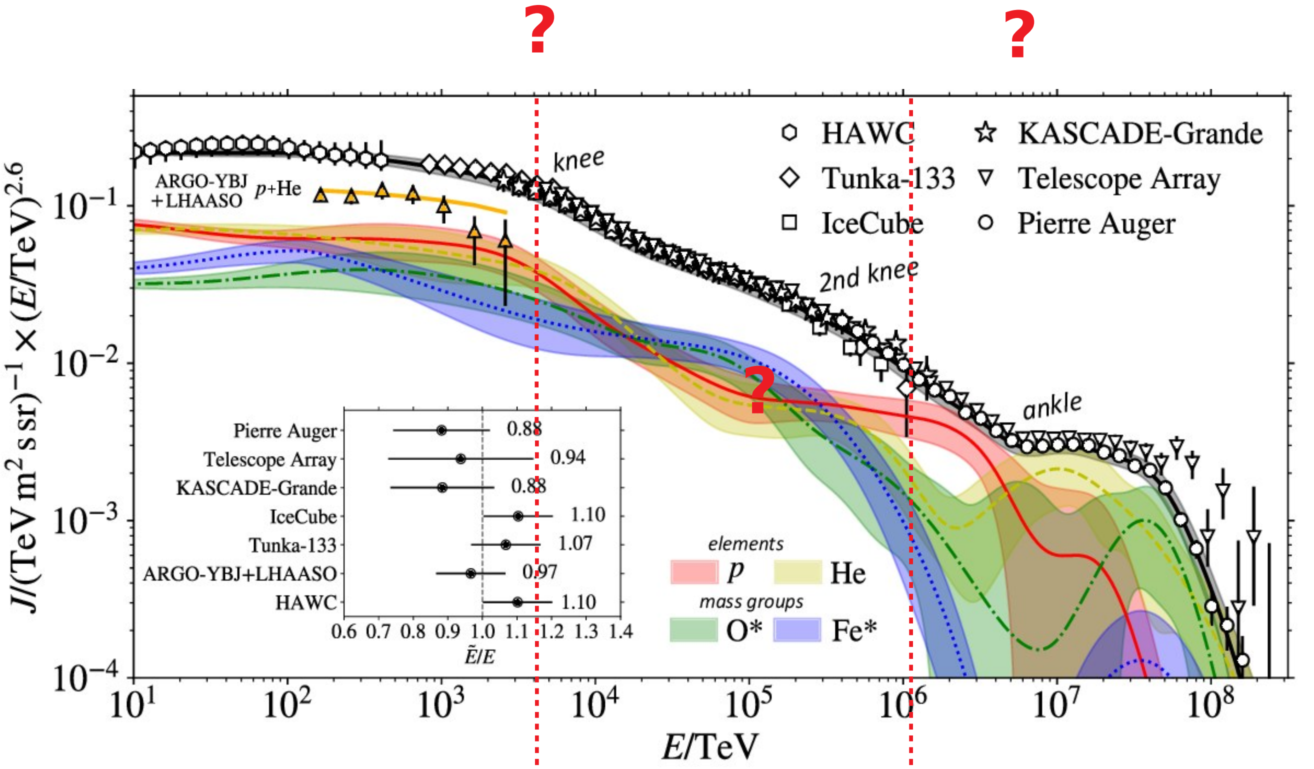
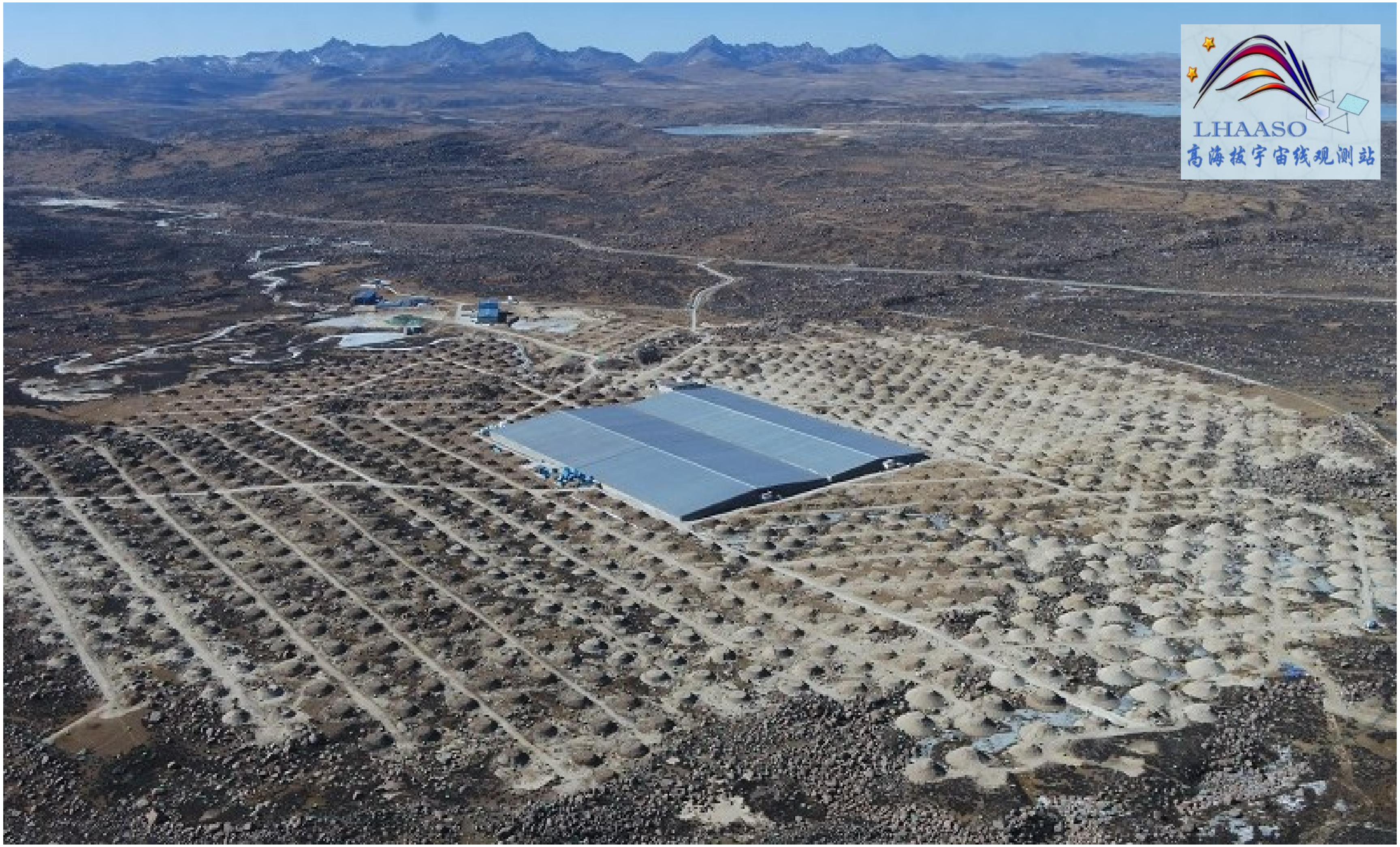
Cosmic-Ray Spectrum



Up to **~10^20 eV**!

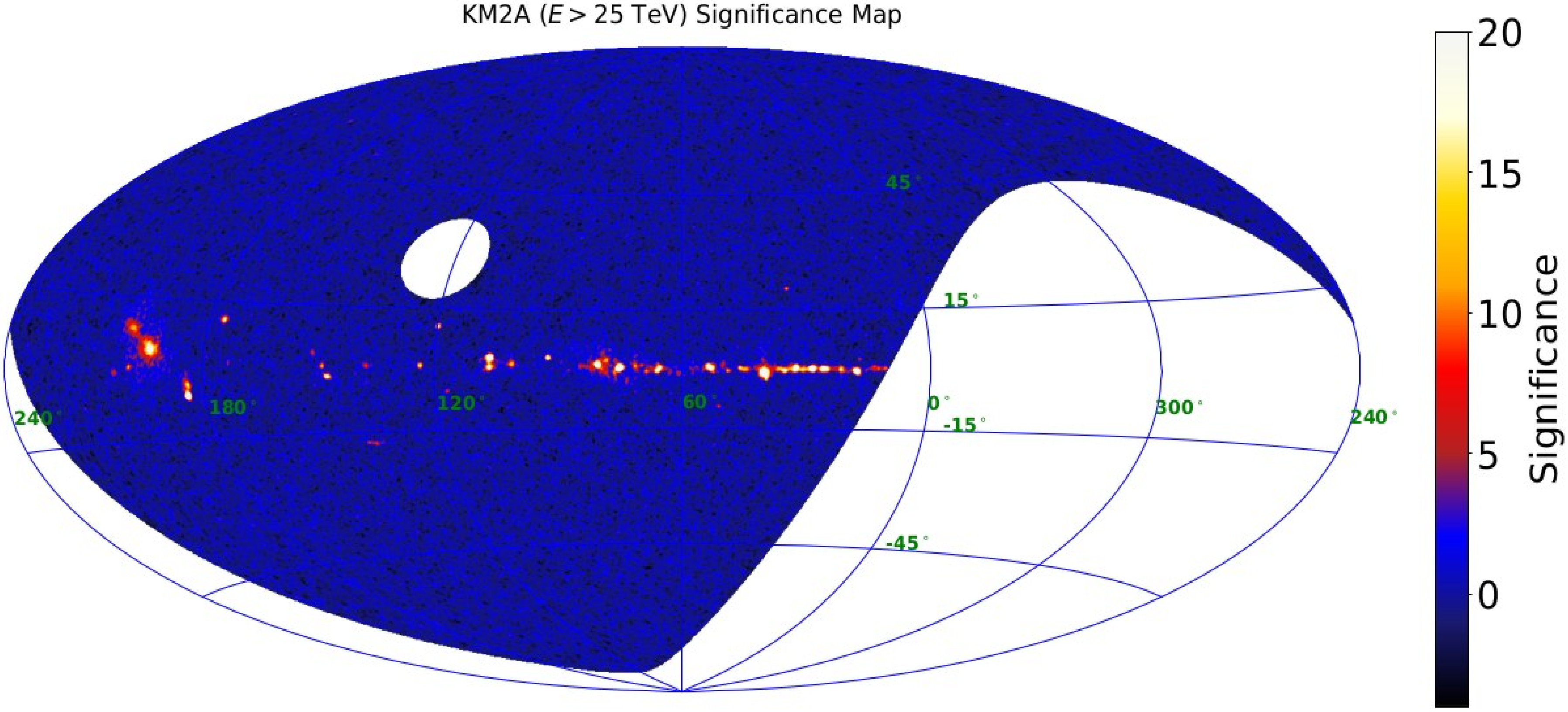
Galactic **TRANSITION** **extragalactic**

PeV Gamma-Ray Astronomy: LHAASO



Sky at 25 TeV – 1 PeV with LHAASO

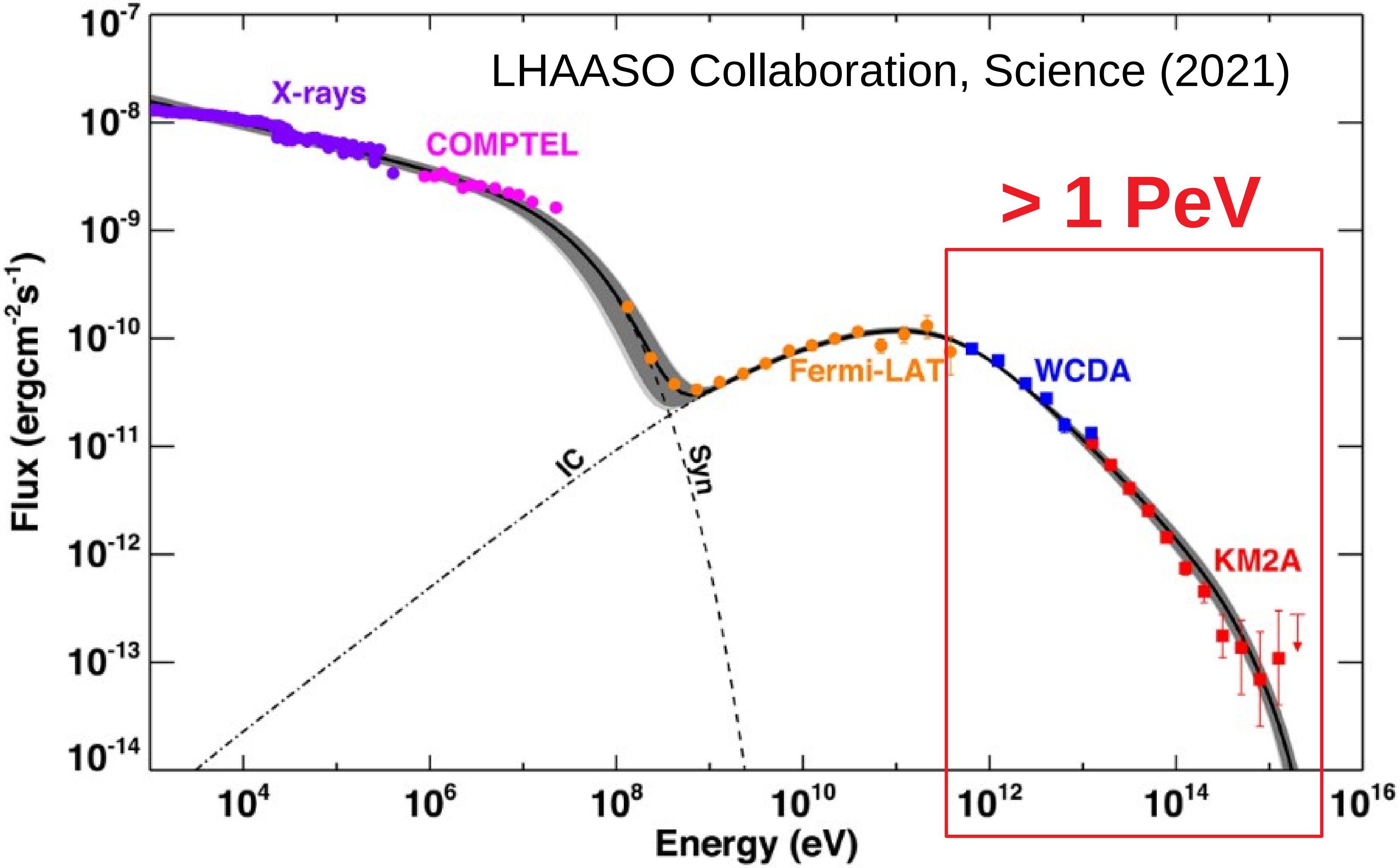
LHAASO Collaboration



> Dozen of PeVatrons:

Many **PWNe and TeV halos** (leptonic) ; **Microquasars** (Hadronic?)

Crab Nebula observed by LHAASO



… and dozens of other PWNe

PWNe & Potential TeV Halos in LHAASO Catalogue

Table 4. 1LHAASO sources associated pulsars

Source name	PSR name	Sep.(°)	d (kpc)	τ_c (kyr)	$\dot{E}$ (erg s^{-1})	P_c	Identified type in TeVCat
1LHAASO J0007+7303u	PSR J0007+7303	0.05	1.40	14	4.5e+35	7.3e-05	PWN
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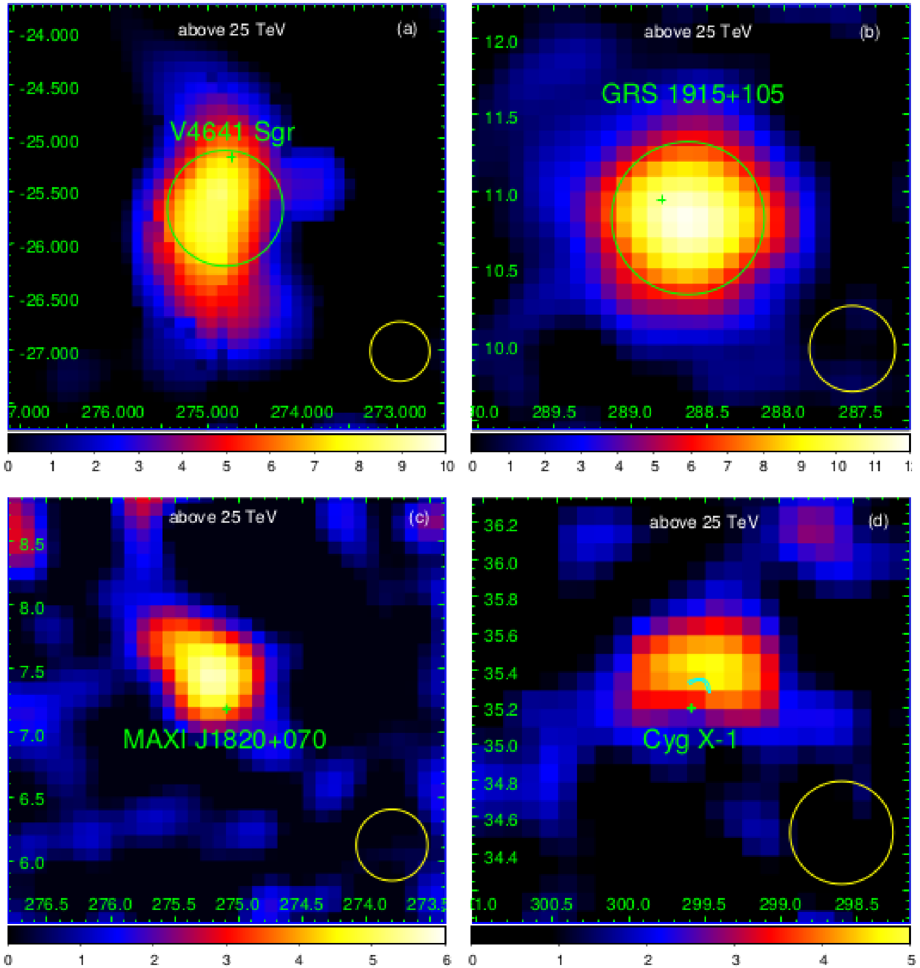
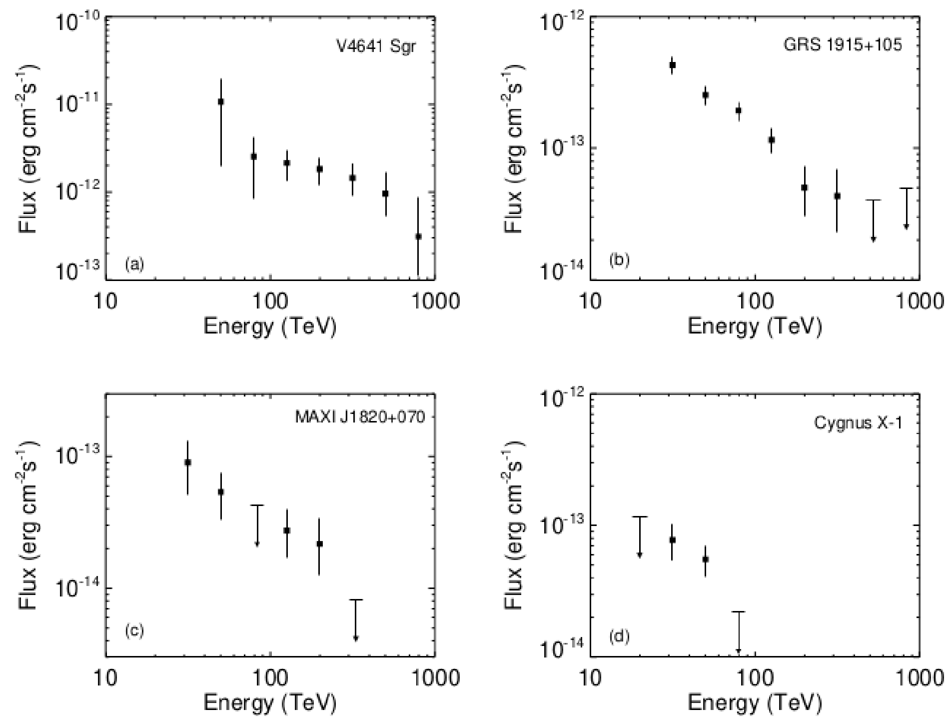
1LHAASO J1740+0948u	PSR J1740+1000	0.21	1.23	114	2.3e+35	1.4e-03	
1LHAASO J1809-1918u	PSR J1809-1917	0.05	3.27	51	1.8e+36	6.2e-04	
1LHAASO J1813-1245	PSR J1813-1245	0.01	2.63	43	6.2e+36	6.3e-06	
1LHAASO J1825-1256u	PSR J1826-1256	0.09	1.55	14	3.6e+36	1.6e-03	
1LHAASO J1825-1337u	PSR J1826-1334	0.11	3.61	21	2.8e+36	2.8e-03	PWN/TeV Halo
1LHAASO J1837-0654u	PSR J1838-0655	0.12	6.60	23	5.6e+36	2.2e-03	PWN
1LHAASO J1839-0548u	PSR J1838-0537	0.20	-	5	6.0e+36	6.1e-03	
1LHAASO J1848-0001u	PSR J1849-0001	0.06	-	43	9.8e+36	1.2e-04	PWN
1LHAASO J1857+0245	PSR J1856+0245	0.16	6.32	21	4.6e+36	3.1e-03	PWN
1LHAASO J1906+0712	PSR J1906+0722	0.19	-	49	1.0e+36	5.9e-03	
1LHAASO J1908+0615u	PSR J1907+0602	0.23	2.37	20	2.8e+36	6.8e-03	
1LHAASO J1912+1014u	PSR J1913+1011	0.13	4.61	169	2.9e+36	1.5e-03	
1LHAASO J1914+1150u	PSR J1915+1150	0.09	14.01	116	5.4e+35	1.8e-03	
1LHAASO J1928+1746u	PSR J1928+1746	0.04	4.34	83	1.6e+36	1.6e-04	
1LHAASO J1929+1846u	PSR J1930+1852	0.29	7.00	3	1.2e+37	2.6e-03	PWN
1LHAASO J1954+2836u	PSR J1954+2836	0.01	1.96	69	1.1e+36	1.6e-05	PWN
1LHAASO J1954+3253	PSR J1952+3252	0.33	3.00	107	3.7e+36	6.7e-03	
1LHAASO J1959+2846u	PSR J1958+2845	0.10	1.95	22	3.4e+35	2.8e-03	PWN
1LHAASO J2005+3415	PSR J2004+3429	0.25	10.78	18	5.8e+35	9.9e-03	
1LHAASO J2005+3050	PSR J2006+3102	0.20	6.04	104	2.2e+35	9.2e-03	
1LHAASO J2020+3649u	PSR J2021+3651	0.05	1.80	17	3.4e+36	1.5e-04	PWN
1LHAASO J2028+3352	PSR J2028+3332	0.36	-	576	3.5e+34	8.0e-03	
1LHAASO J2031+4127u	PSR J2032+4127	0.08	1.33	201	1.5e+35	1.0e-03	PWN
1LHAASO J2228+6100u	PSR J2229+6114	0.27	3.00	10	2.2e+37	2.2e-03	PWN
1LHAASO J2238+5900	PSR J2238+5903	0.07	2.83	27	8.9e+35	3.0e-04	

LHAASO Collaboration, ApJS 271, 25 (2024)

LHAASO's PeVatron microquasars

LHAASO Collaboration, arXiv:2410.08988





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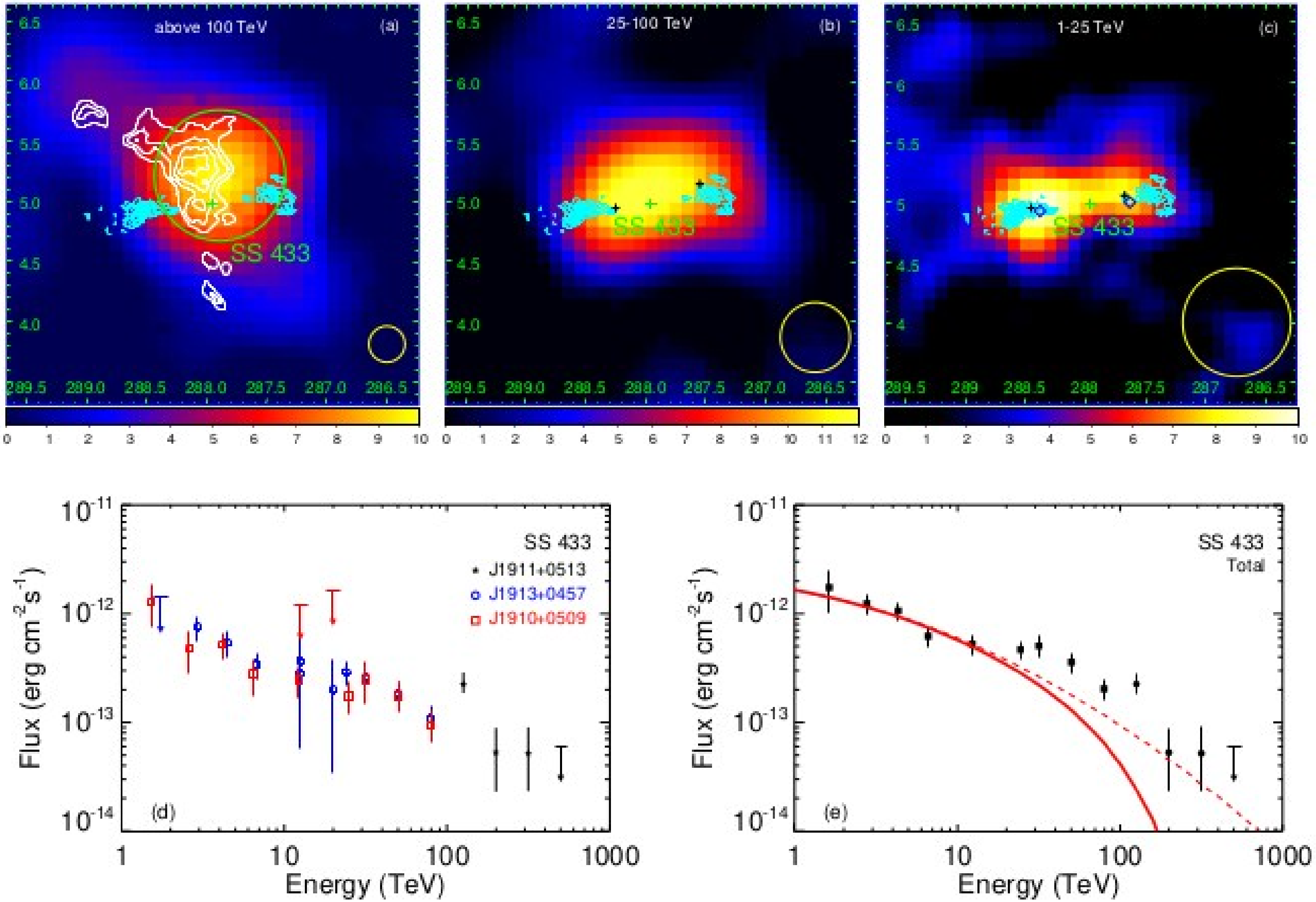
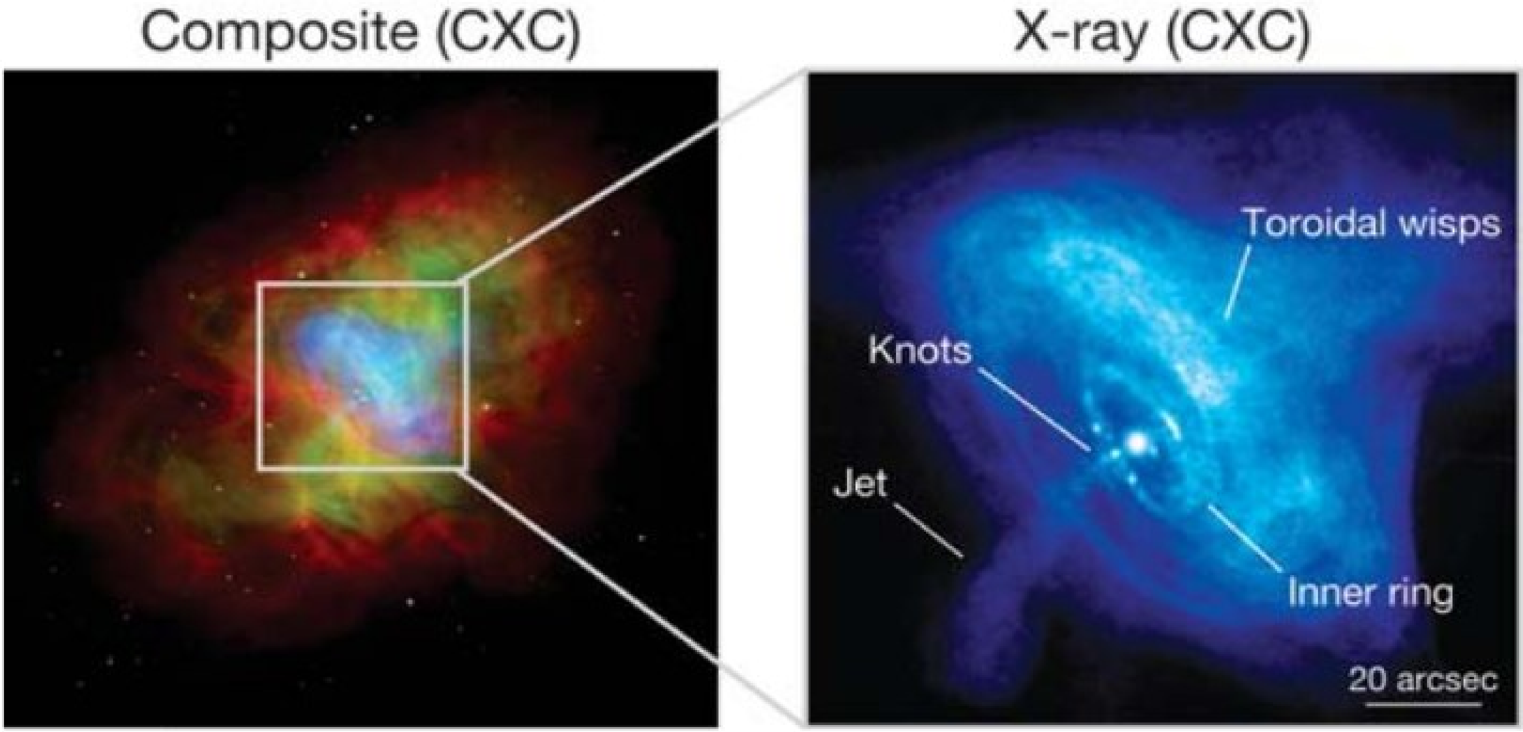


1 – Pulsar Wind Nebulae & TeV halos

Synchrotron from the Crab Nebula

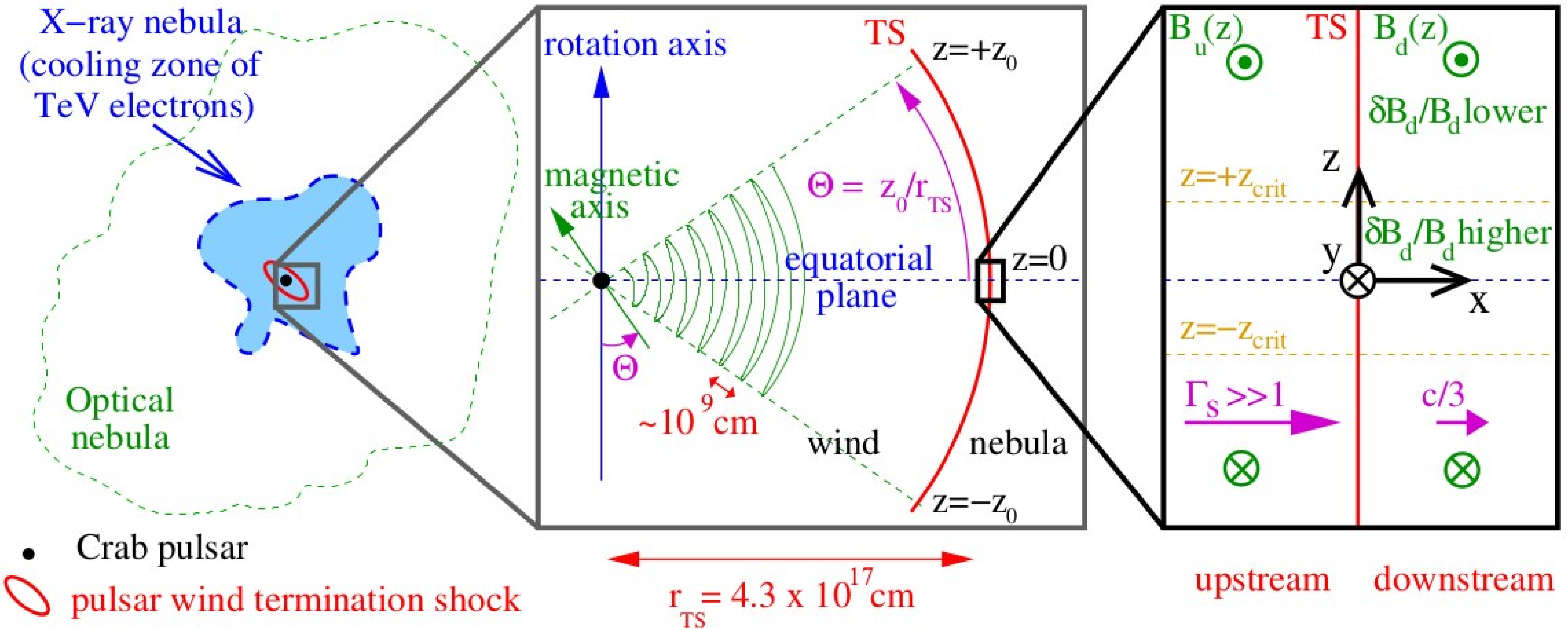


BUT perpendicular shock

=> 1st order Fermi should NOT work !

Model

Giacinti & Kirk, ApJ **863**, 18 (2018), arXiv:1804.05056



→ **INJECTION :** e.g.
Giacchè & Kirk (2017):

$$E_{\rm inj,d} = 1\,{\rm TeV}$$

→ Komissarov profile for the B field,

→ Add synthetic turbulence,

→ Integrate the **trajectories** of **individual particules** in **3D** (**test particle limit**).

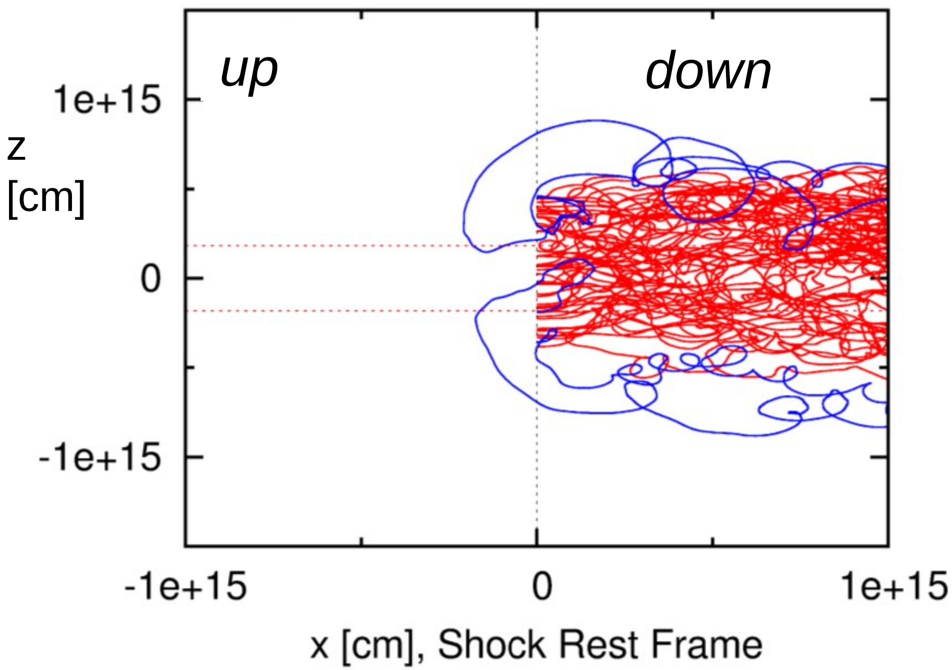
Numerical simulations

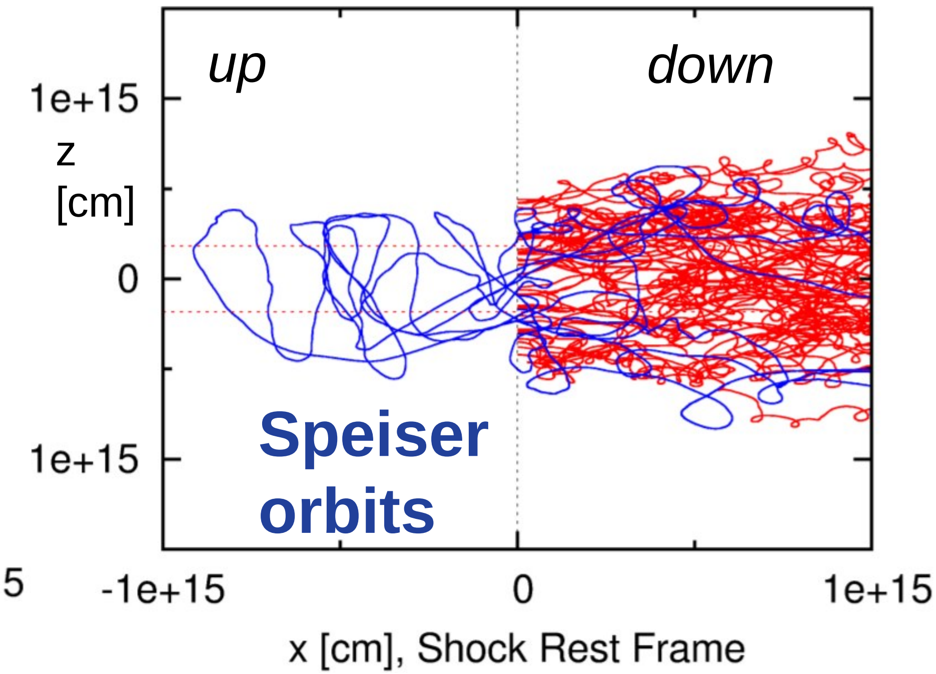
<u>POSITRONS:</u> **<u>ELECTRONS:</u>**

(roles exchanged with opposite pulsar polarity)

No/little acceleration

Efficient acceleration to VHE



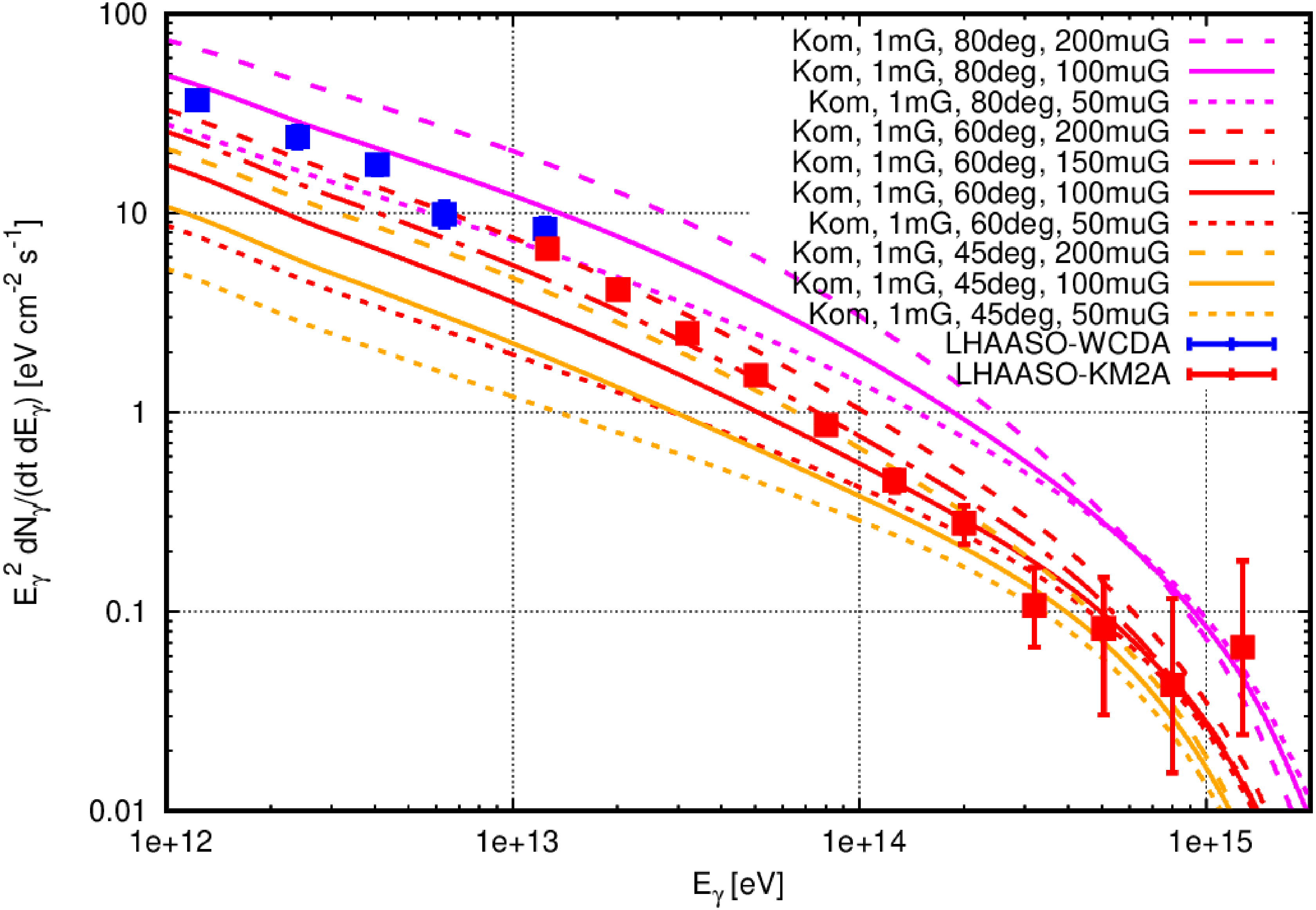


Drift motions on the shock surface eject the positrons...

...and **confine the electrons in the equatorial region**

PeV γ-rays from the Crab Nebula

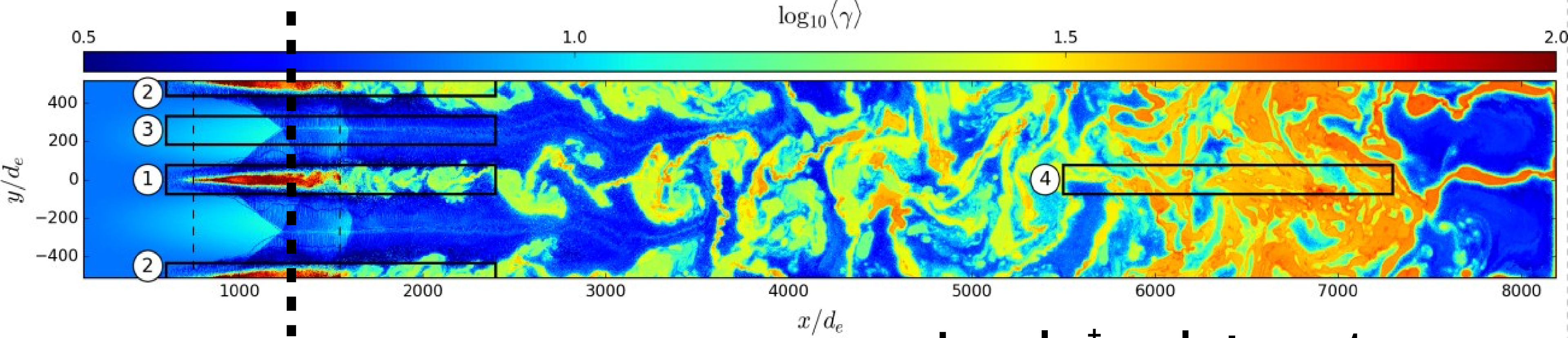
Comparison with LHAASO data :



Particle-In-Cell simulations

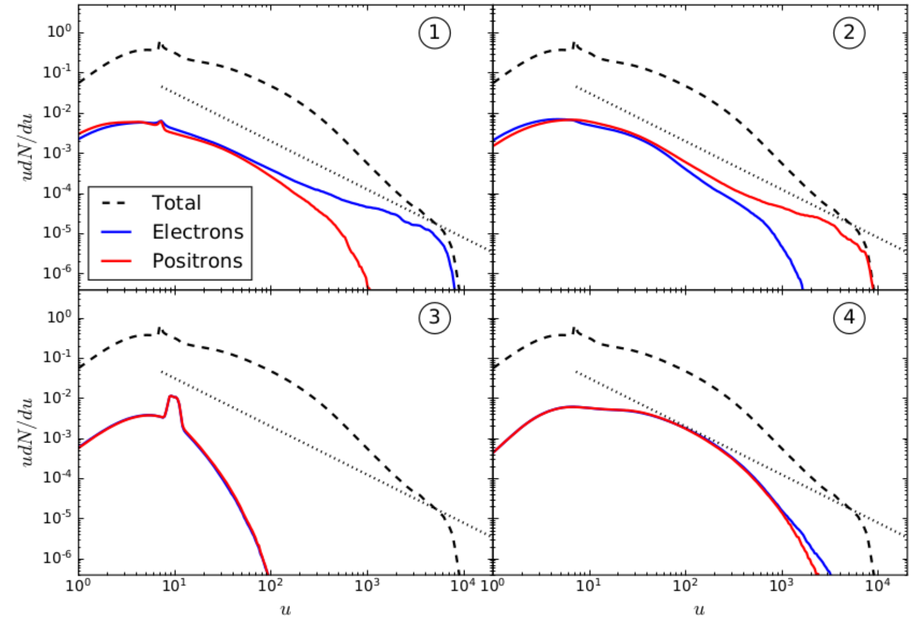
Cerutti & Giacinti, A&A **642**, A123 (2020)

Mean particle Lorentz factor at $\omega_{pe} t = 7840$ for $\sigma_0 = 30$:

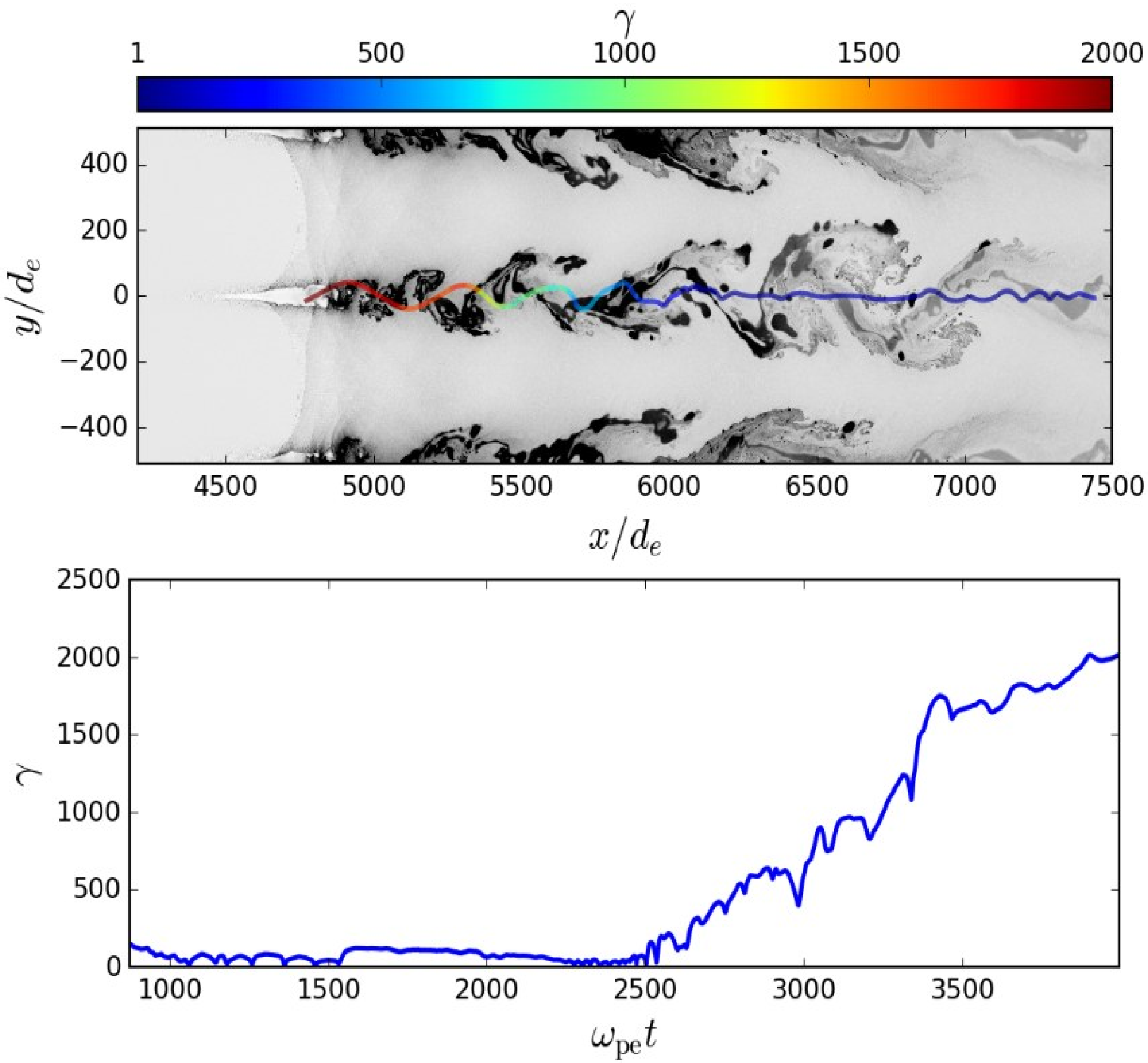


Shock

Local e^+ and e^- spectra :



Particle trajectories / acc. mechanisms:

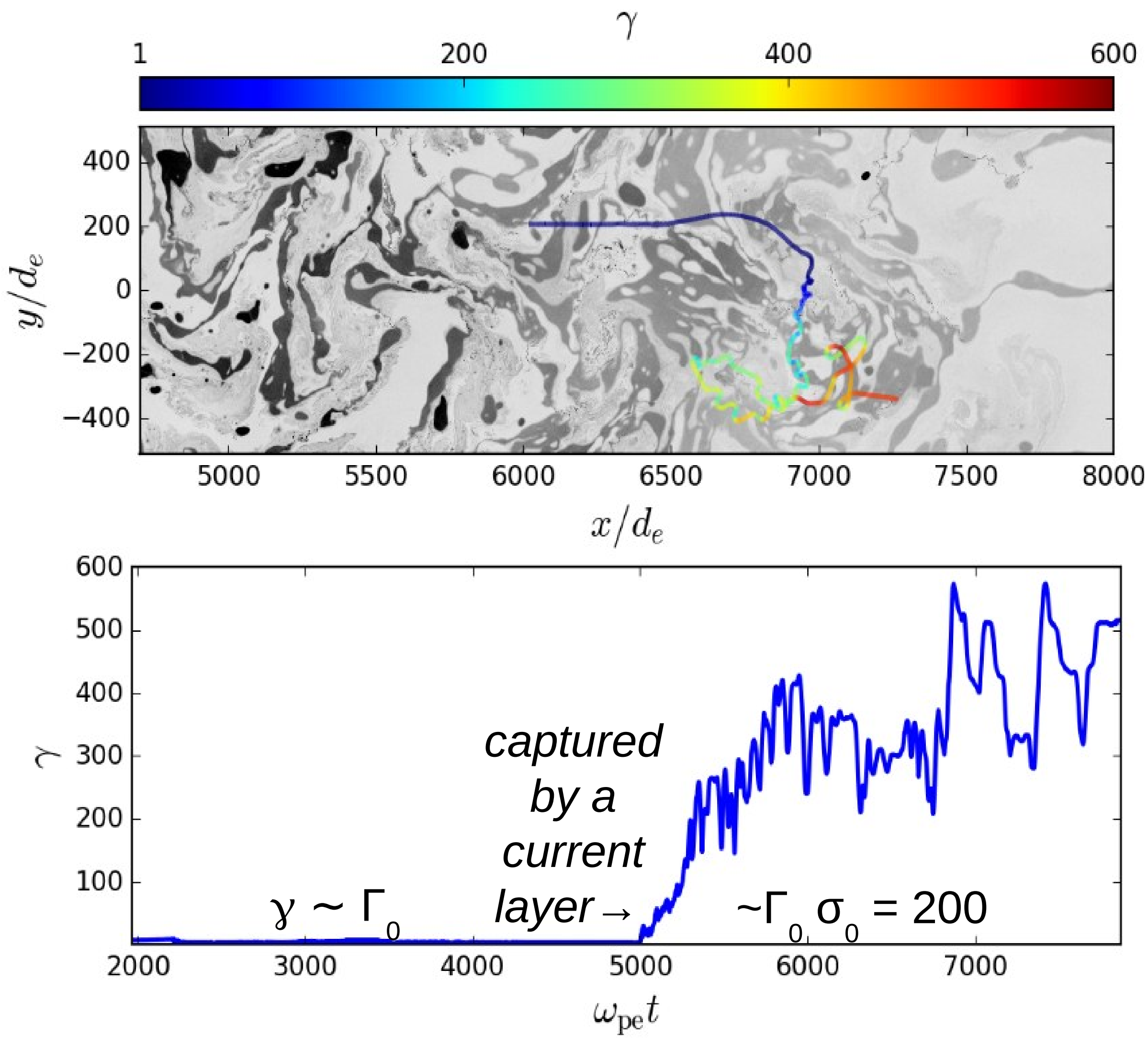


→ Particles on **Speiser orbits** in the **equatorial cavity**

→ **Shear-flow acceleration**

→ Variable, Doppler-boosted synchrotron emission → **Crab flares?**

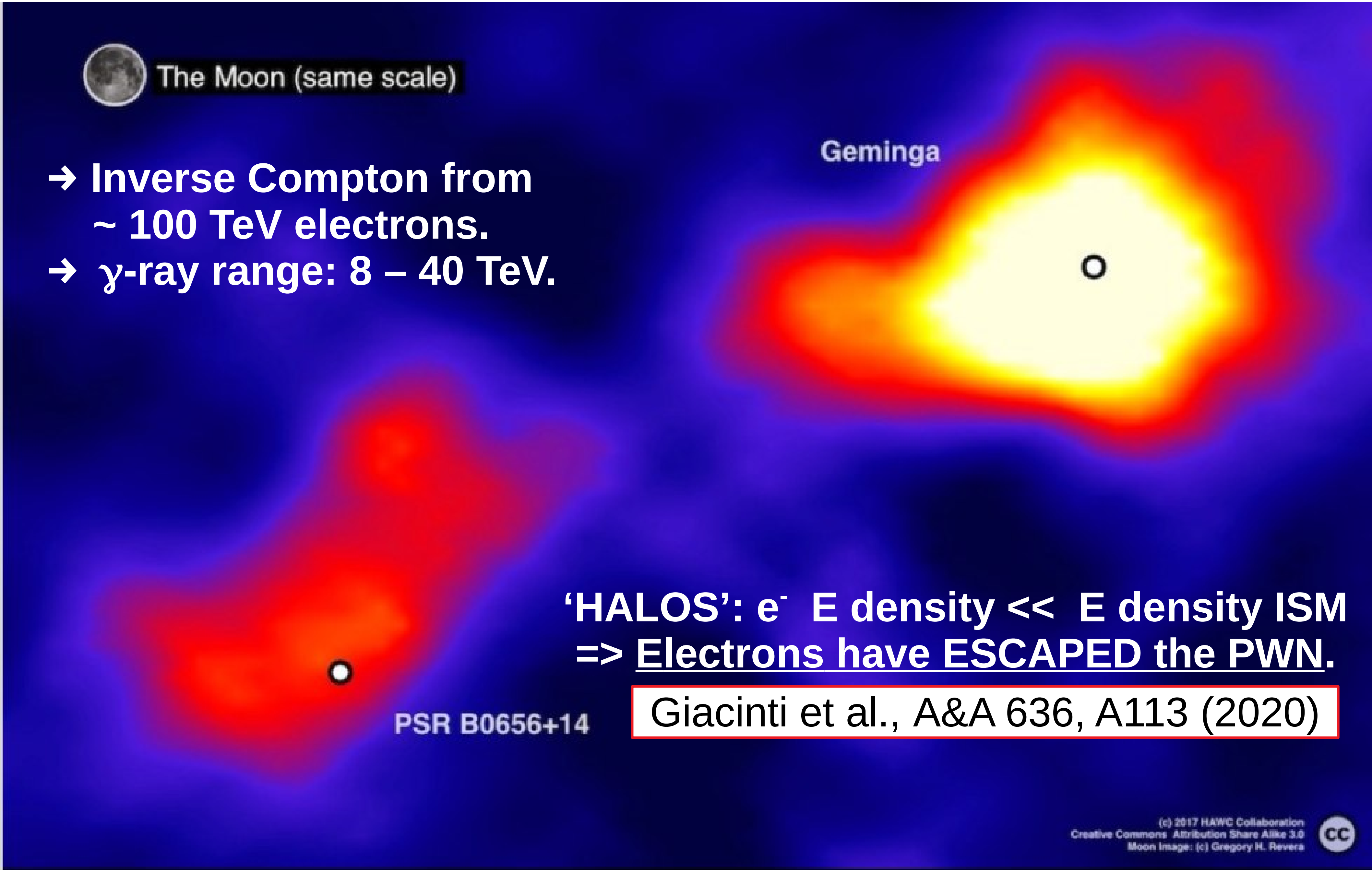
Particle trajectories / acc. mechanisms:



→ Particle captured by a current layer, => **acceleration via rela-tivistic reconnection.**

→ Then (slower) **stochastic acceleration**.

HAWC observ. of Geminga & Monogem



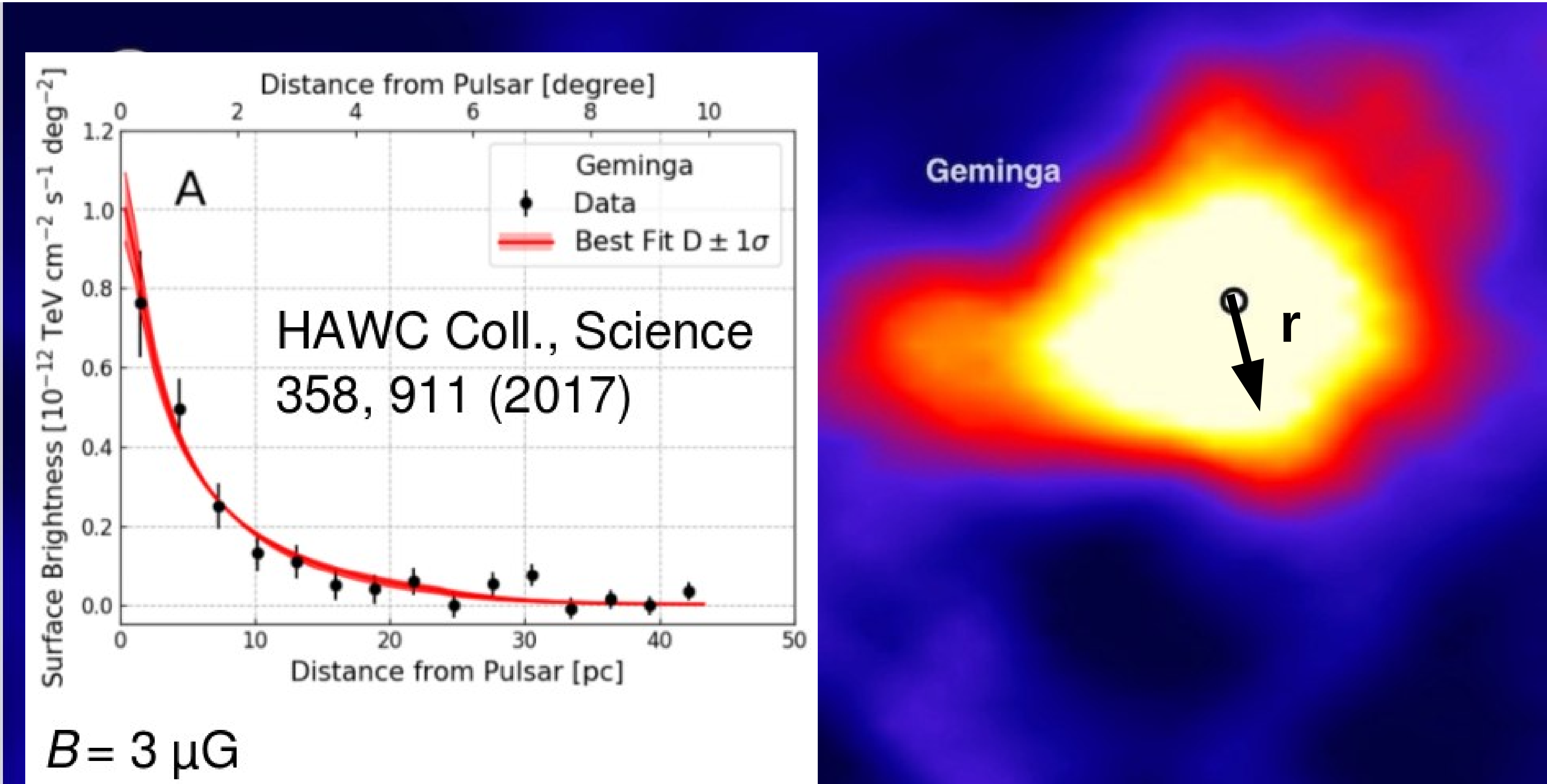
- ➜ Inverse Compton from ~ 100 TeV electrons.
- ➜ γ-ray range: 8 – 40 TeV.

‘HALOS’: e^- E density << E density ISM

=> <u>Electrons have ESCAPED the PWN</u>.

Giacinti et al., A&A 636, A113 (2020)

TeV Halos as probes of ISMF properties



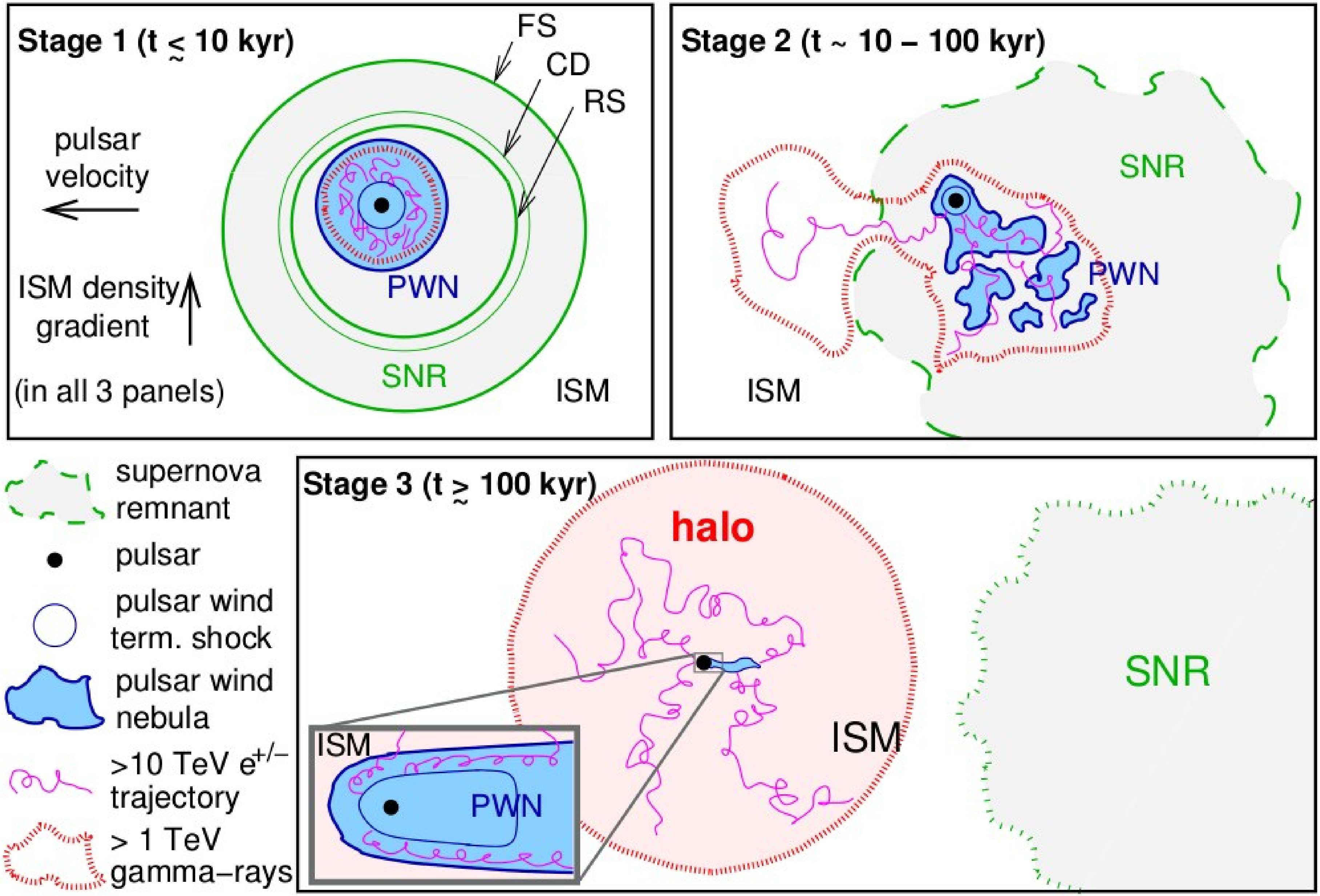
$B = 3\ \mu G$

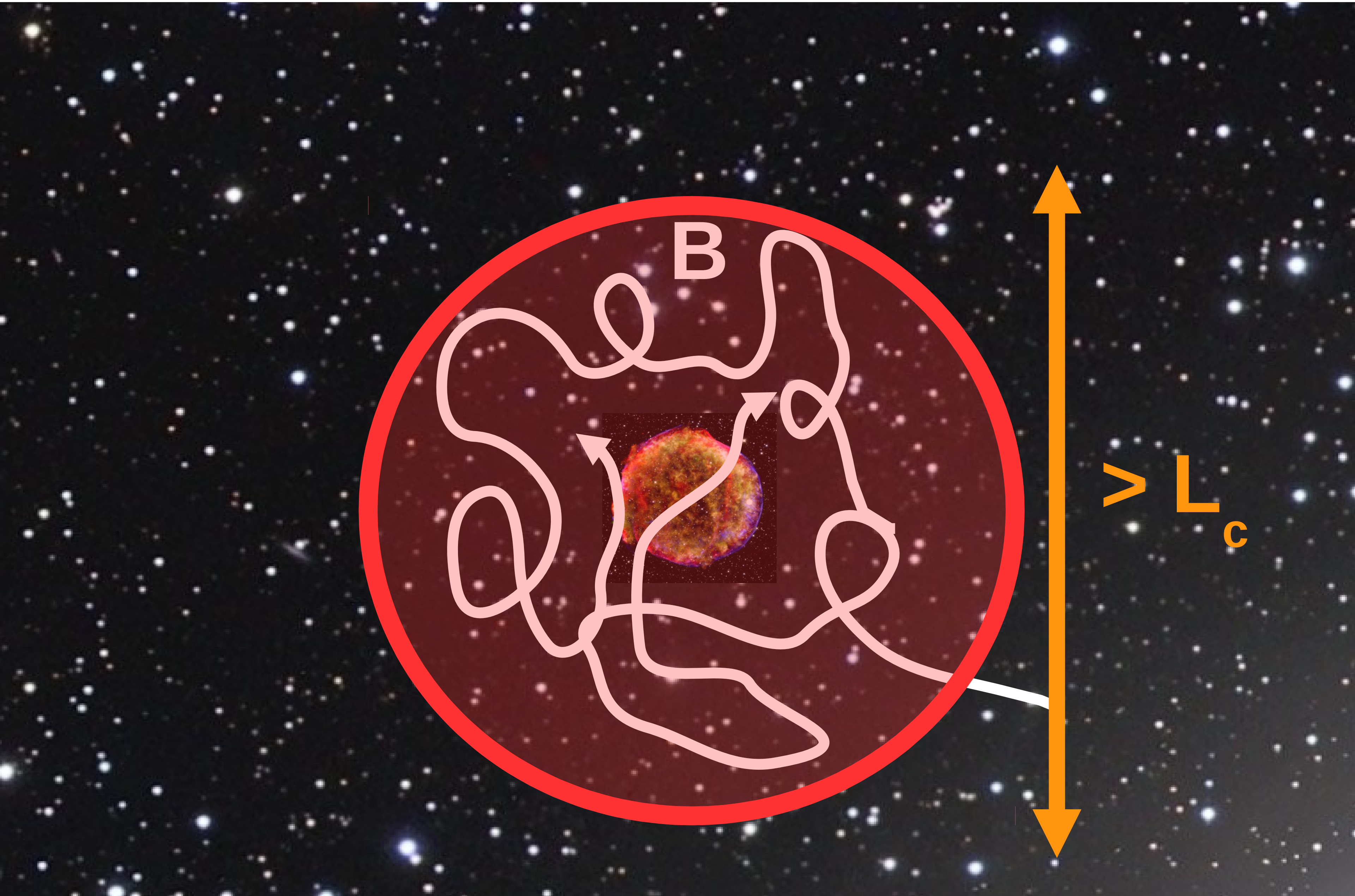
Best fit: **D(100 TeV) = 4.5 x 10^{27} cm^2/s**

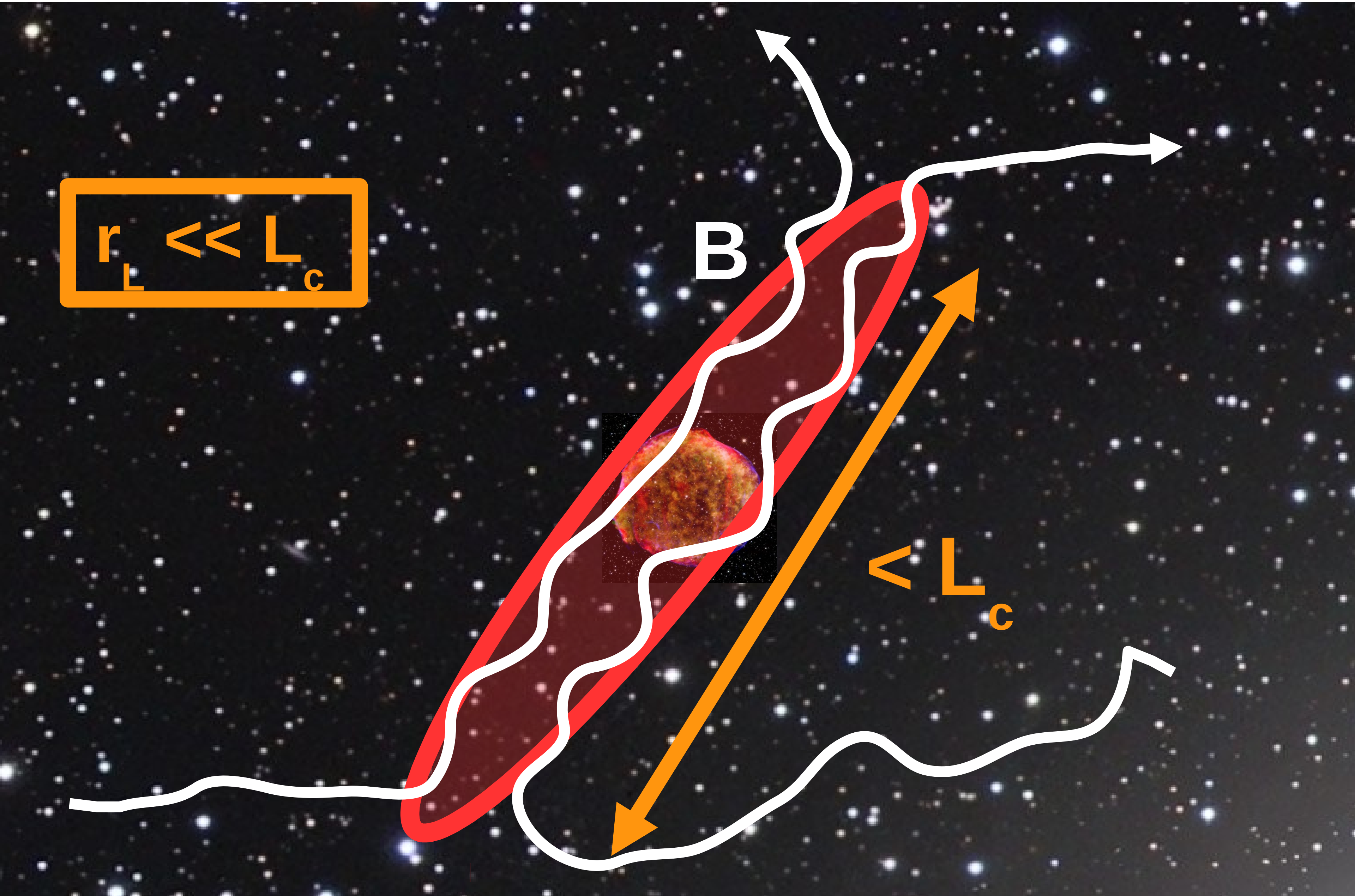
2 orders of magnitude smaller than value from B/C ratio

Evolutionary stages of a PWN :

Giacinti, Mitchell, Lopez-Coto, Joshi, Parsons & Hinton, A&A 636, A113 (2020), arXiv:1907.12121:



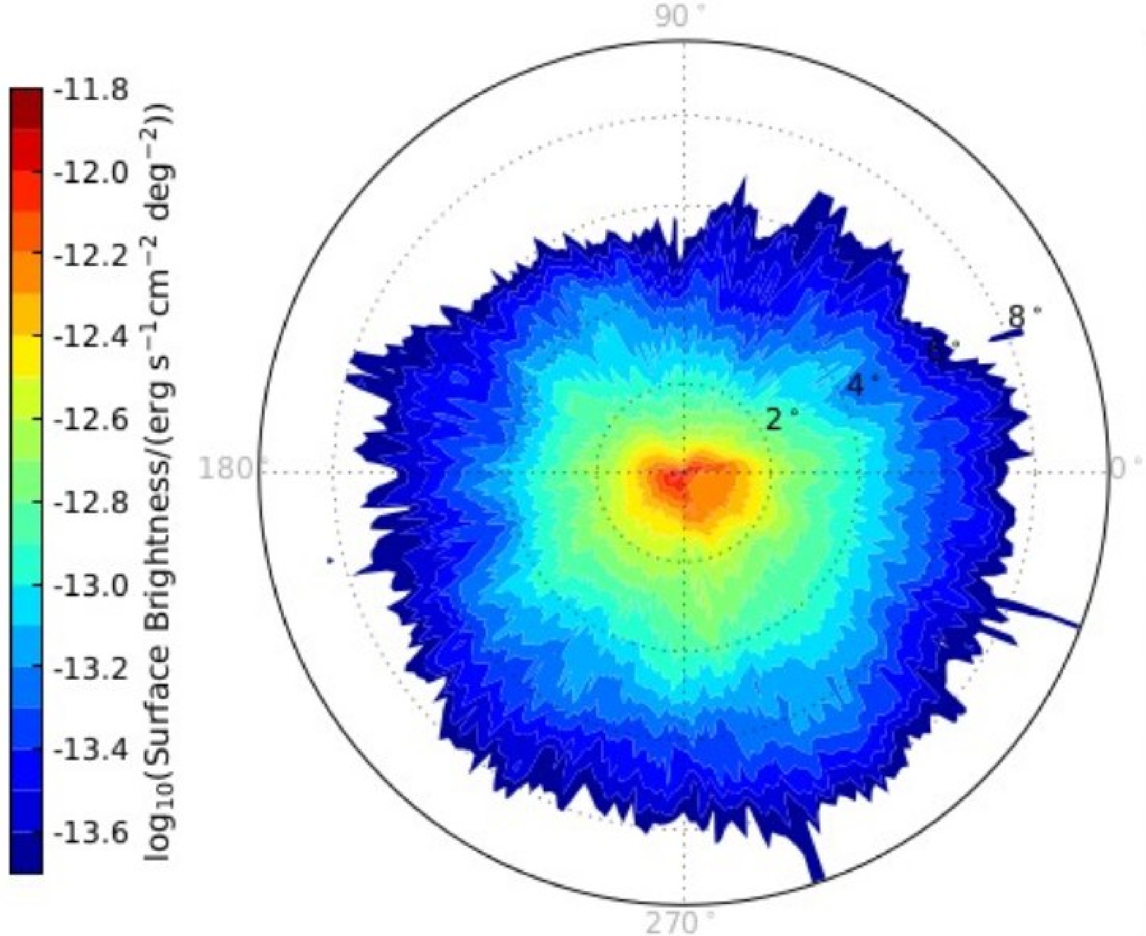

B
> L
c



$r_L << L_c$
B
$< L_c$

Predicted γ-ray surface brightness

Kolmogorov, B_{rms}=3 μG, $\mathbf{L_c}$ **= 5 pc** :

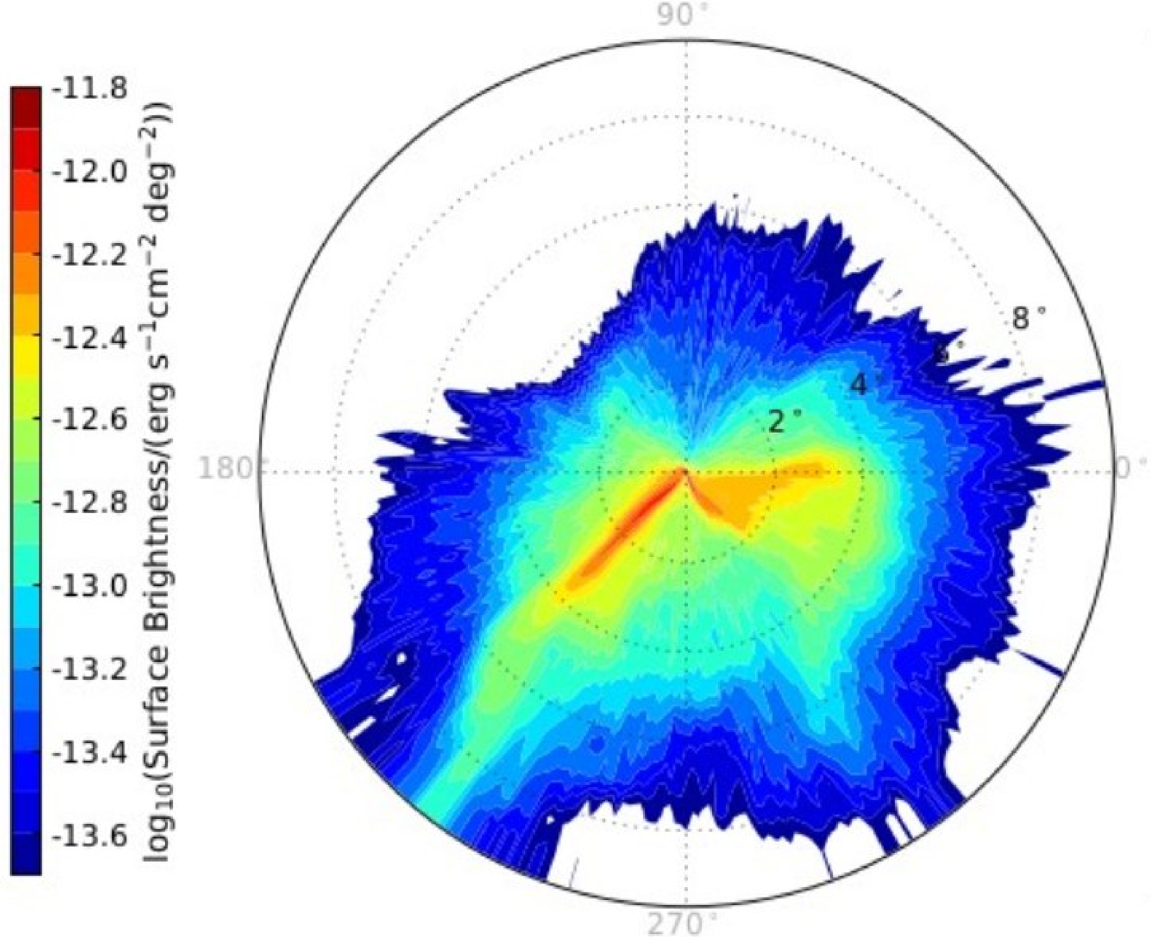


OK

Lopez-Coto & Giacinti, MNRAS 479, 4526 (2018) [arXiv:1712.04373]

Predicted γ-ray surface brightness

Kolmogorov, B_{rms}=3 μG, $\mathbf{L_c}$ **= 40 pc** :

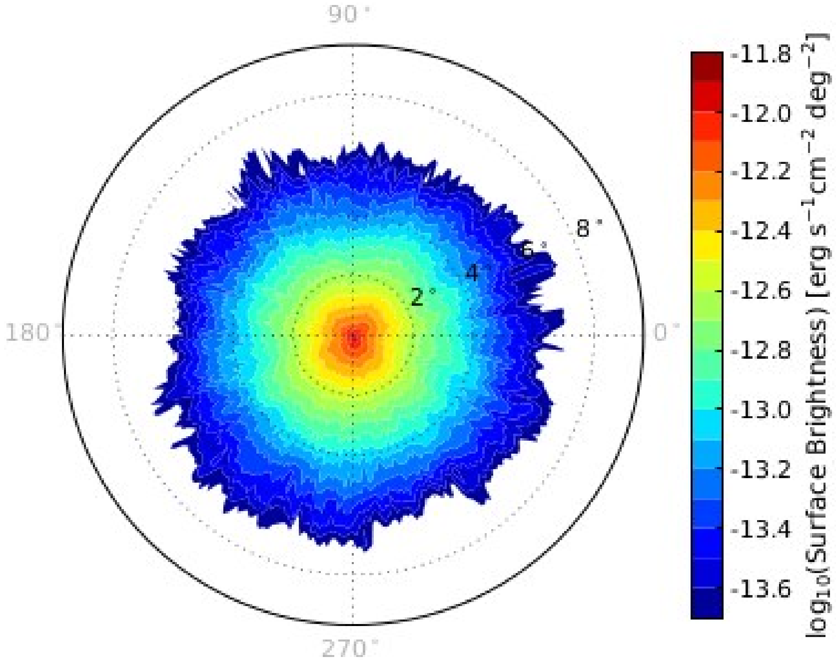


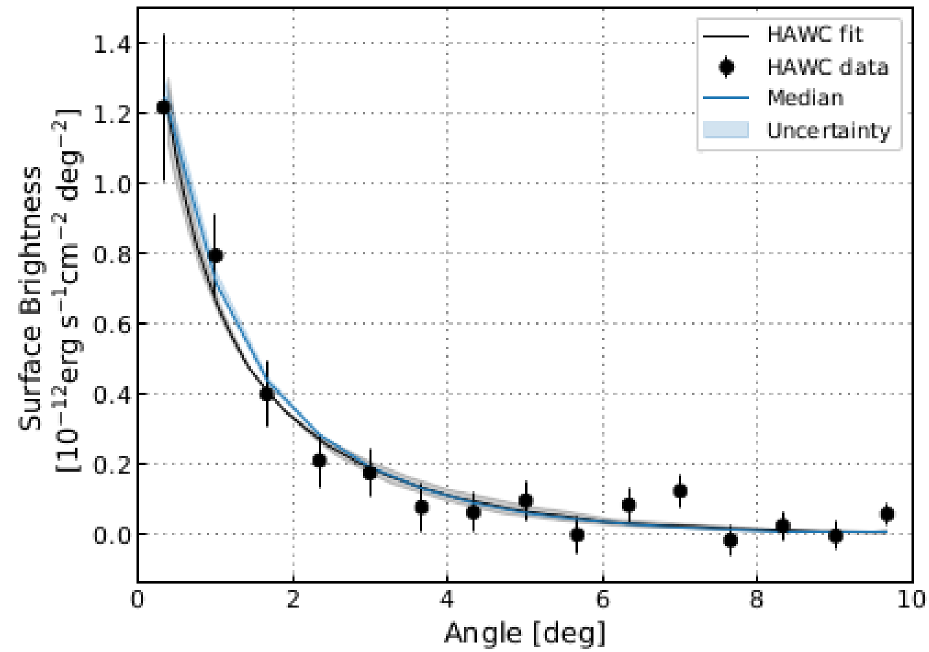
INCOMPATIBLE WITH HAWC MEASUREMENTS

Large coherence lengths (> 10 pc) ruled out (Too asymmetric)

Best fit to HAWC measurements (χ^2/ndf<1)

→ Kolmogorov / Kraichnan, B = 3 μG, L_c = 1 pc





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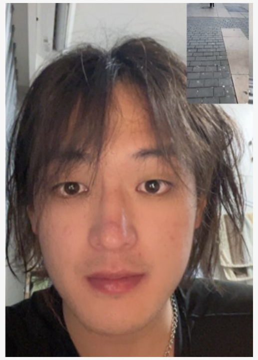
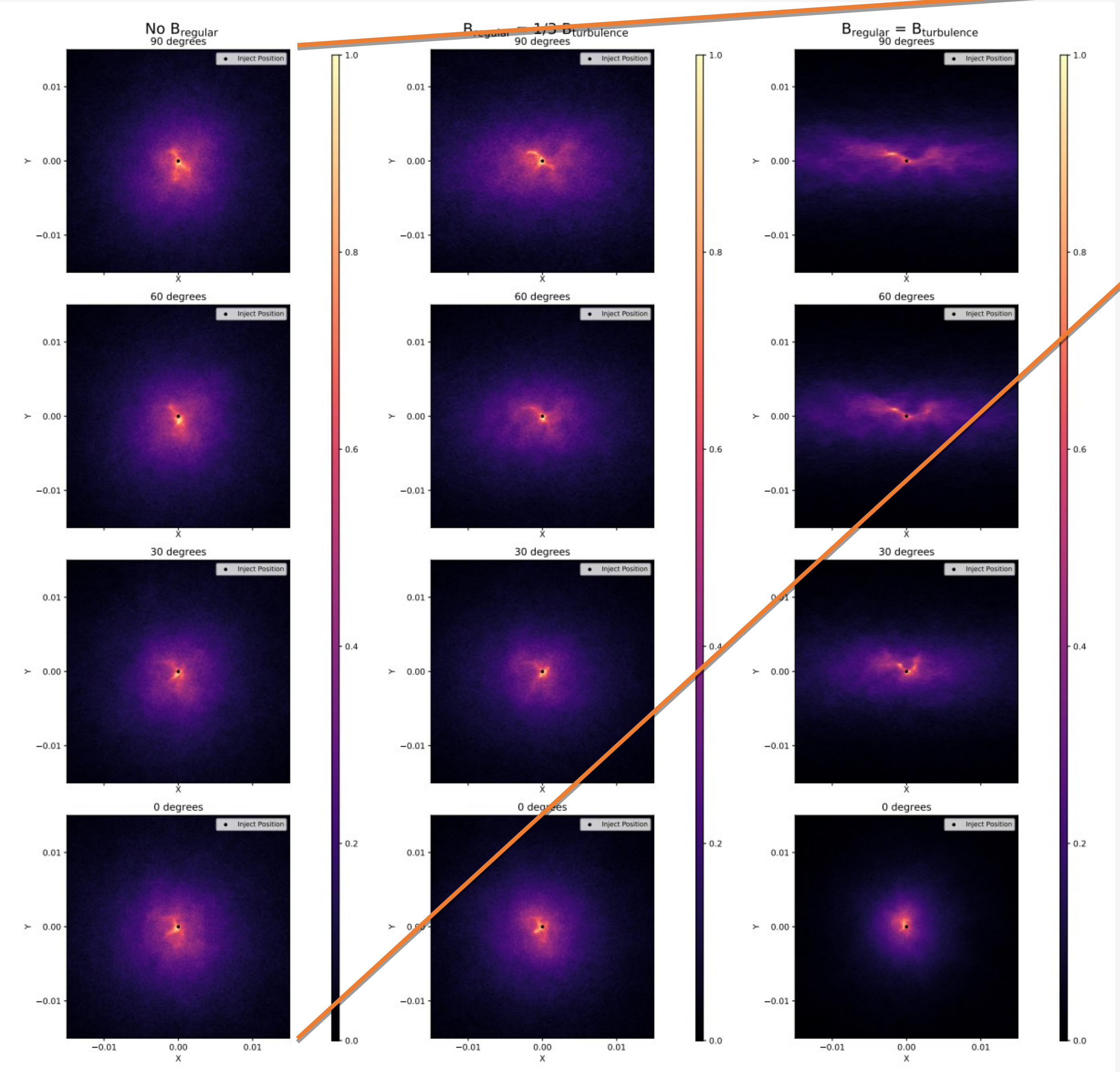
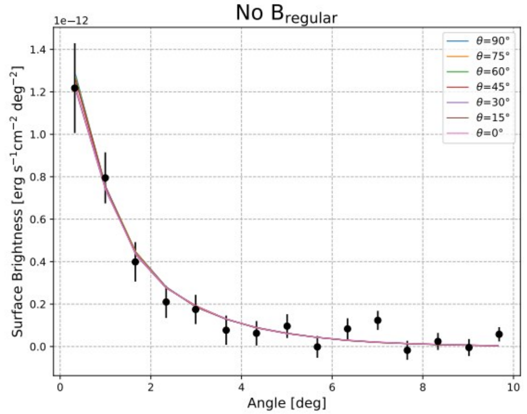
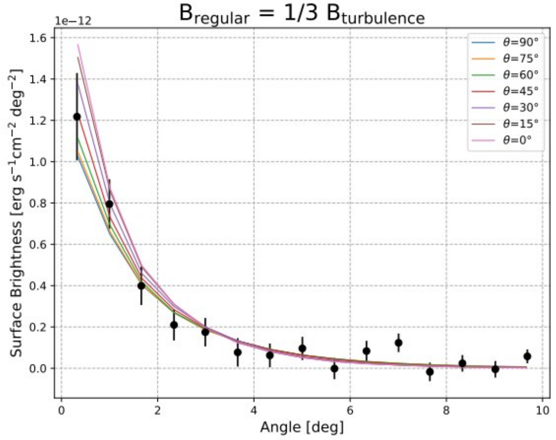
Li, Giacinti & Lopez-Coto, To be submitted (2025)

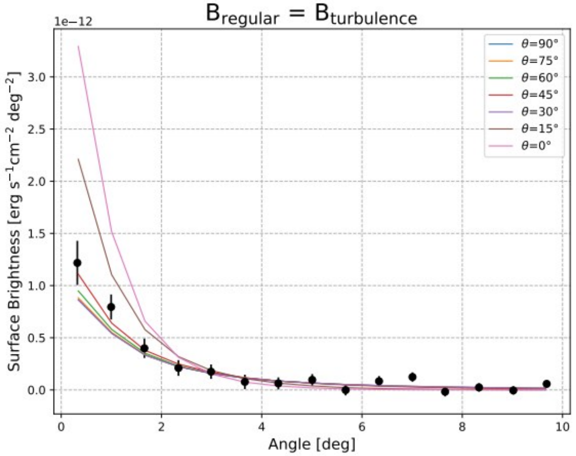


Yuan Li







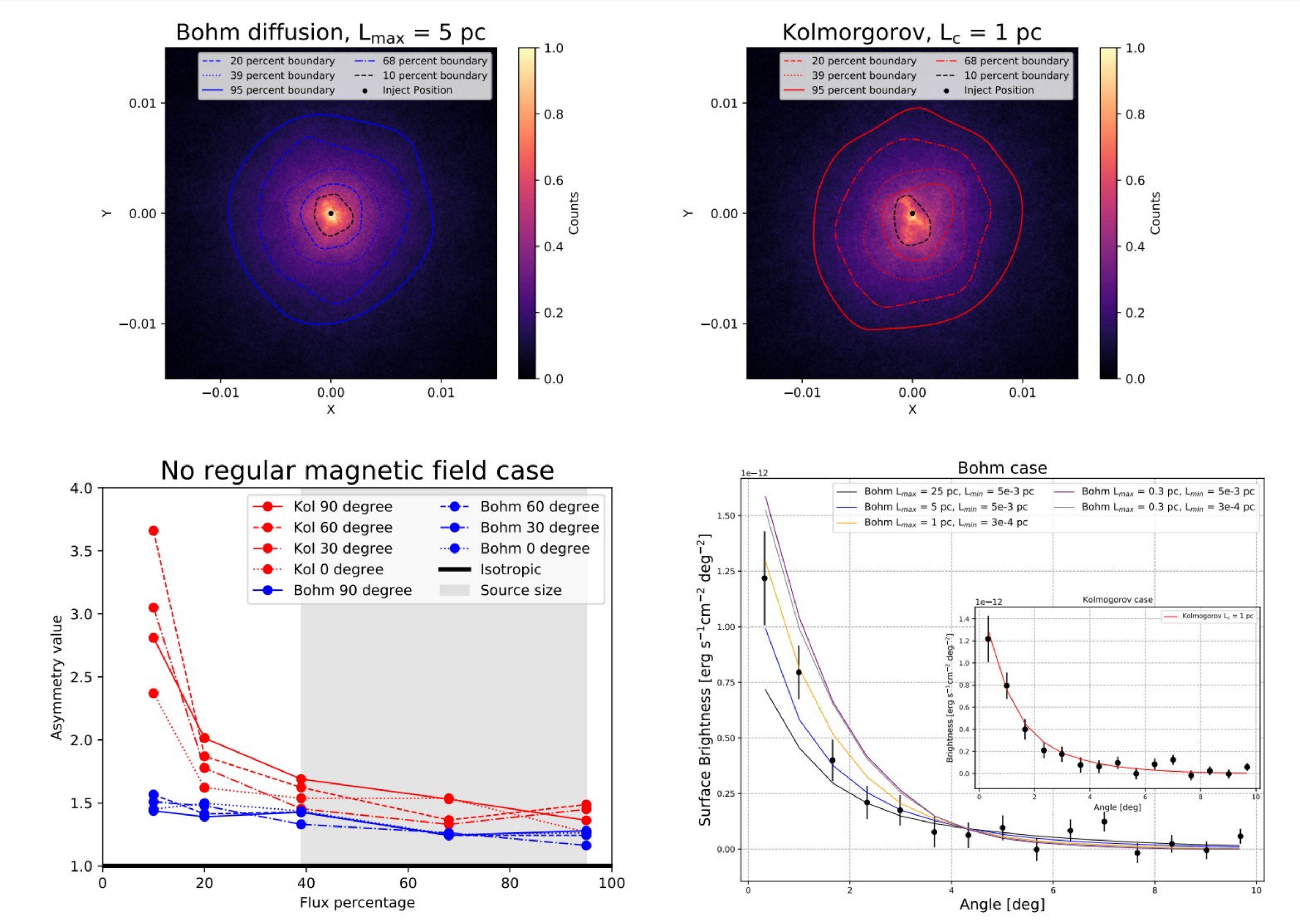


Best fits around
L_c=1-5 pc,
B_{rms}=3-4 μG

Kolmogorov,
L_c=1 pc,
B_{rms}=3 μG

Asymmetry results ($B_{regular} = 0$):

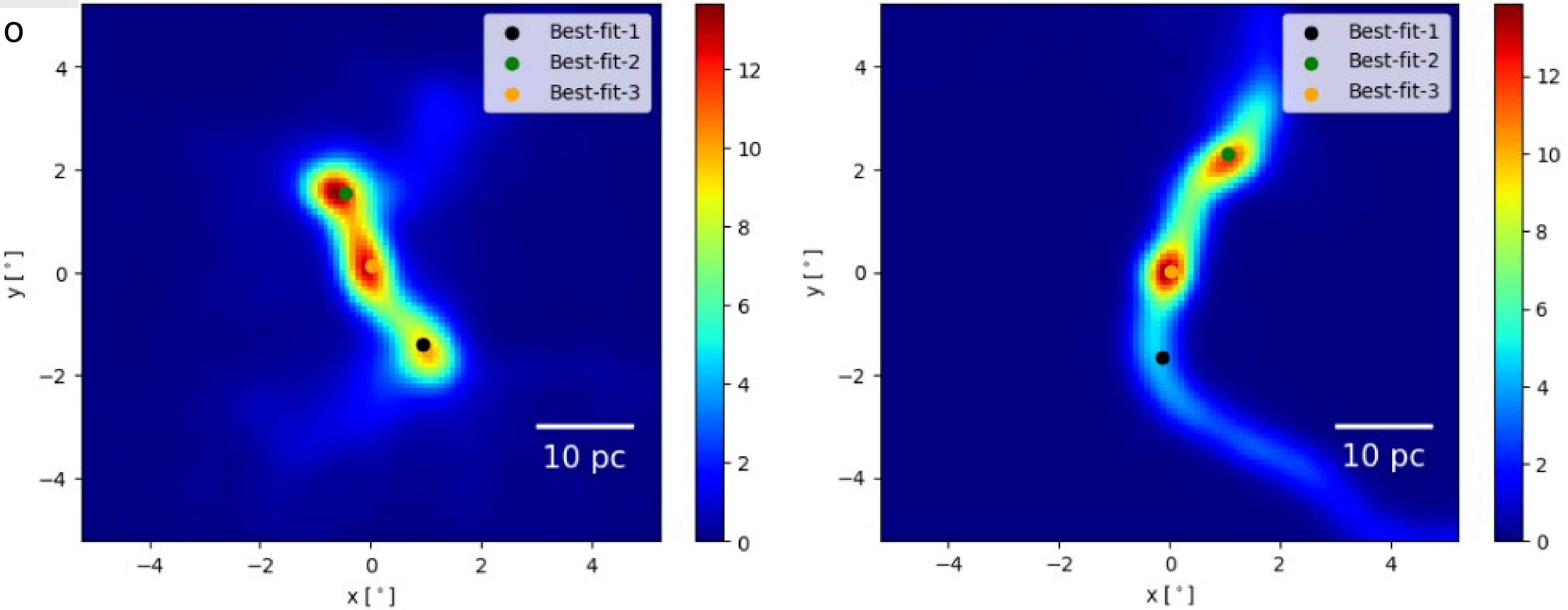
Li, Giacinti & Lopez-Coto, To be submitted (2025)



Yiwei Bao

Appearance of "mirage" sources:

They may appear around astrophysical sources.



L_c = 40pc ; B_{turb} = 3 μG ; B_{reg} = 0 μG; Kolmogorov turbulence ; (8192 particles)

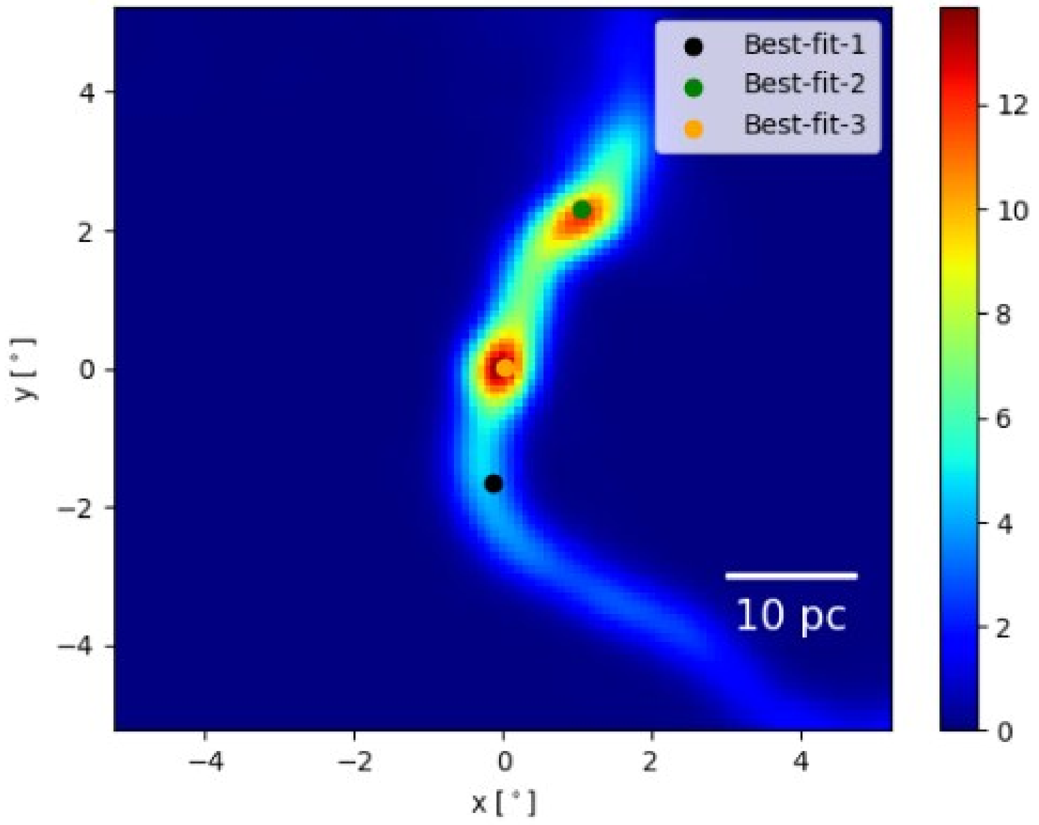
Bao, Giacinti, Liu, Zhang & Chen, arXiv:2407.02478
Bao, Liu, Giacinti, Zhang & Chen, arXiv:2407.02829

Appearance of “mirage” sources:

Yiwei Bao

The second source is a "**mirage**", where the magnetic field bends inwards /outwards, wrt/ observer.

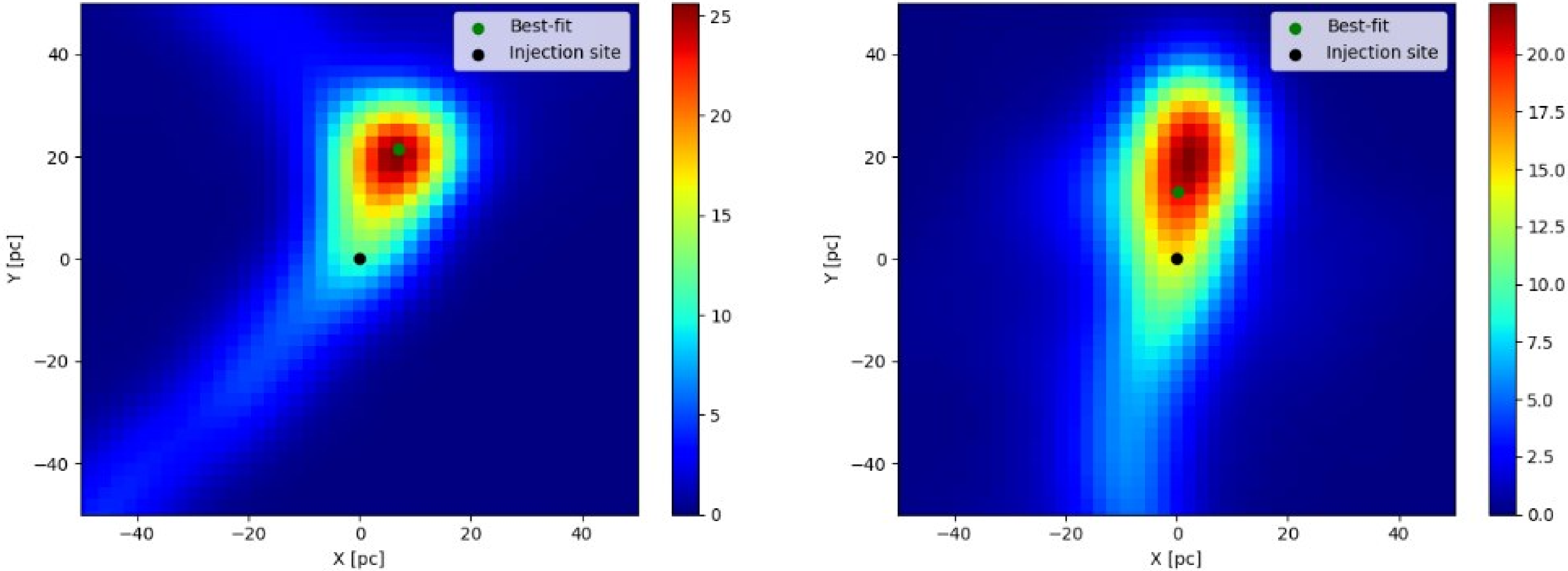
(Prediction: X-ray emission at the mirage source fainter than that at the connecting structure.)



Bao, Giacinti, Liu, Zhang & Chen, arXiv:2407.02478
Bao, Liu, Giacinti, Zhang & Chen, arXiv:2407.02829

Large offsets:

Large offsets may exist between real source and detected source

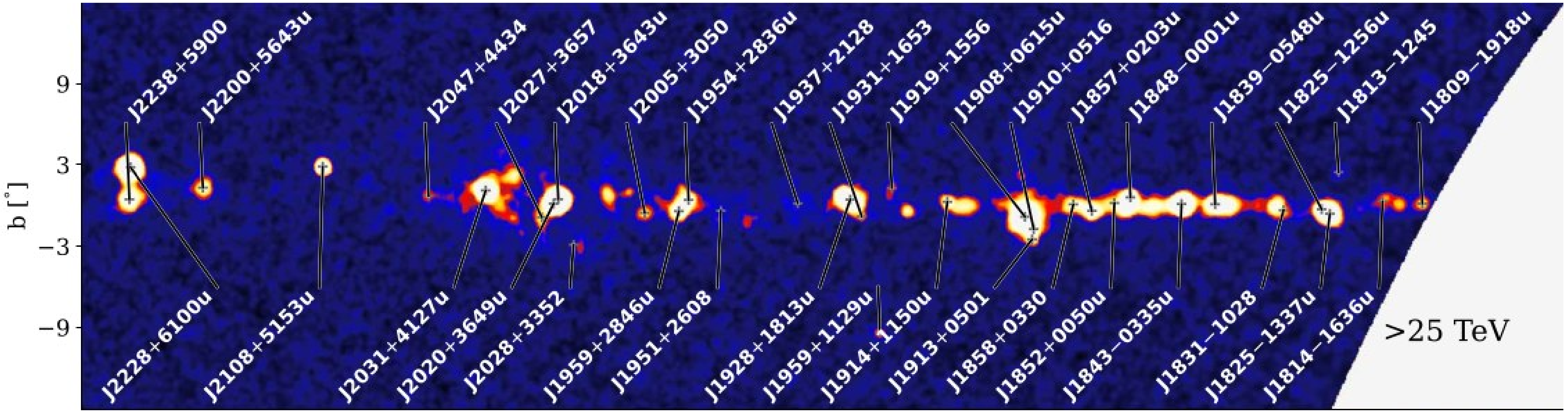


B_{turb} ~ 1 μG ; B_{reg} = 0 μG ; L_c = 200 pc ; Kolmogorov turbulence ; (8192 particles)

Could explain LHAASO observations

LHAASO Collaboration, ApJS 271, 25 (2024)

Many **extended sources** w/ **irregular shapes:**



Large offsets between sources and center

No counterparts?

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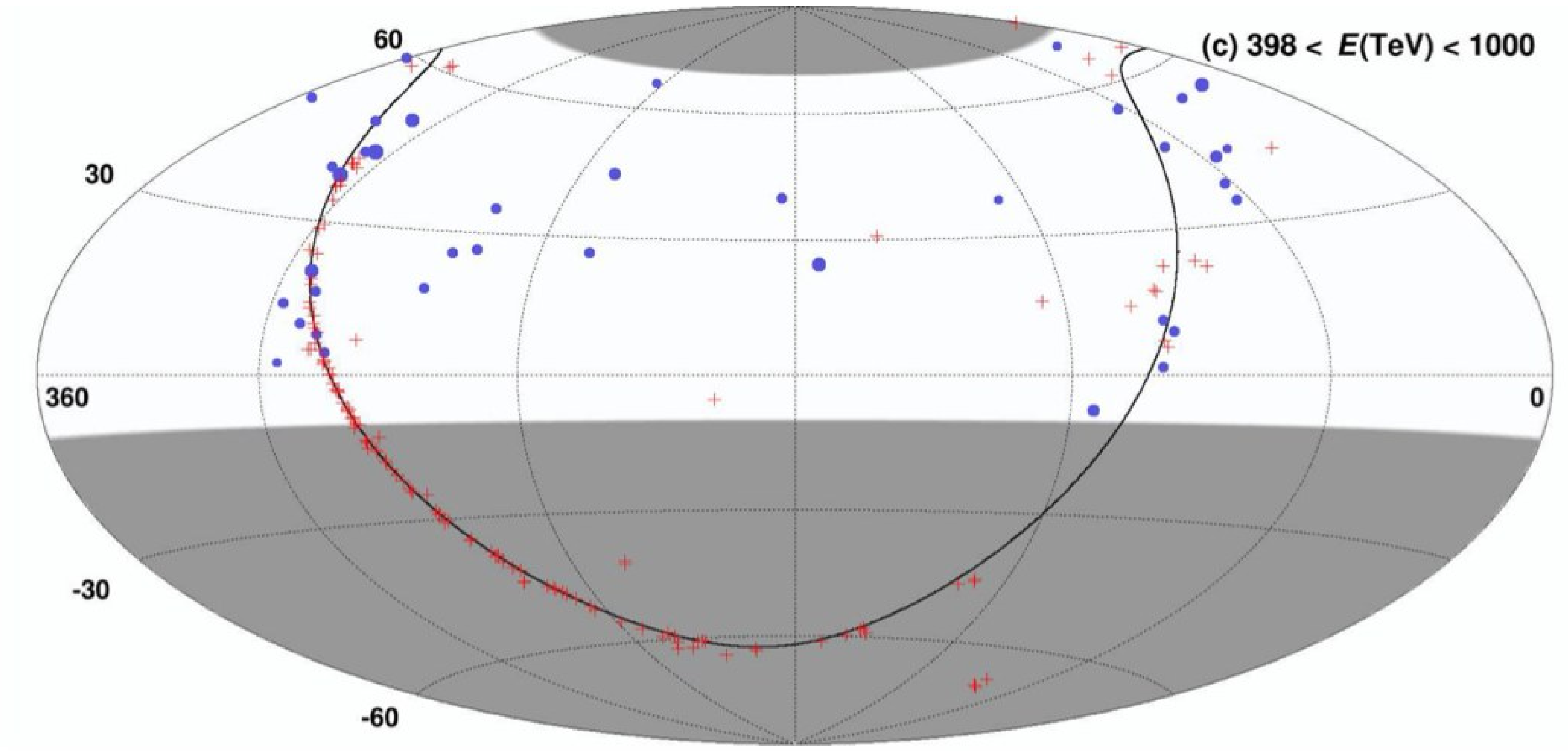
2 – Diffuse γ-ray emission: Pulsar wind nebulae & microquasars

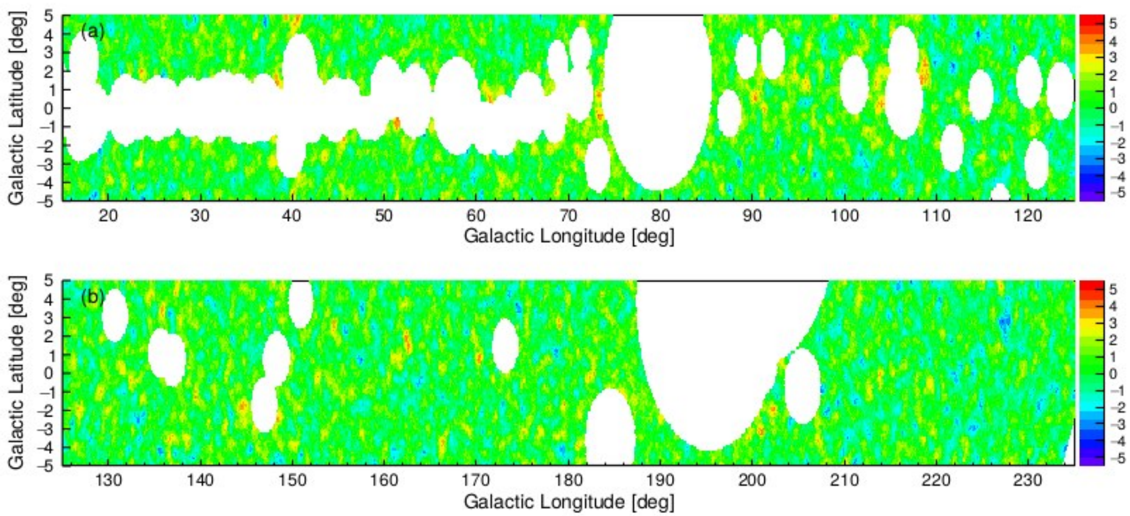
The sky at ~ 400 TeV – 1 PeV: Diffuse emission from AS-γ

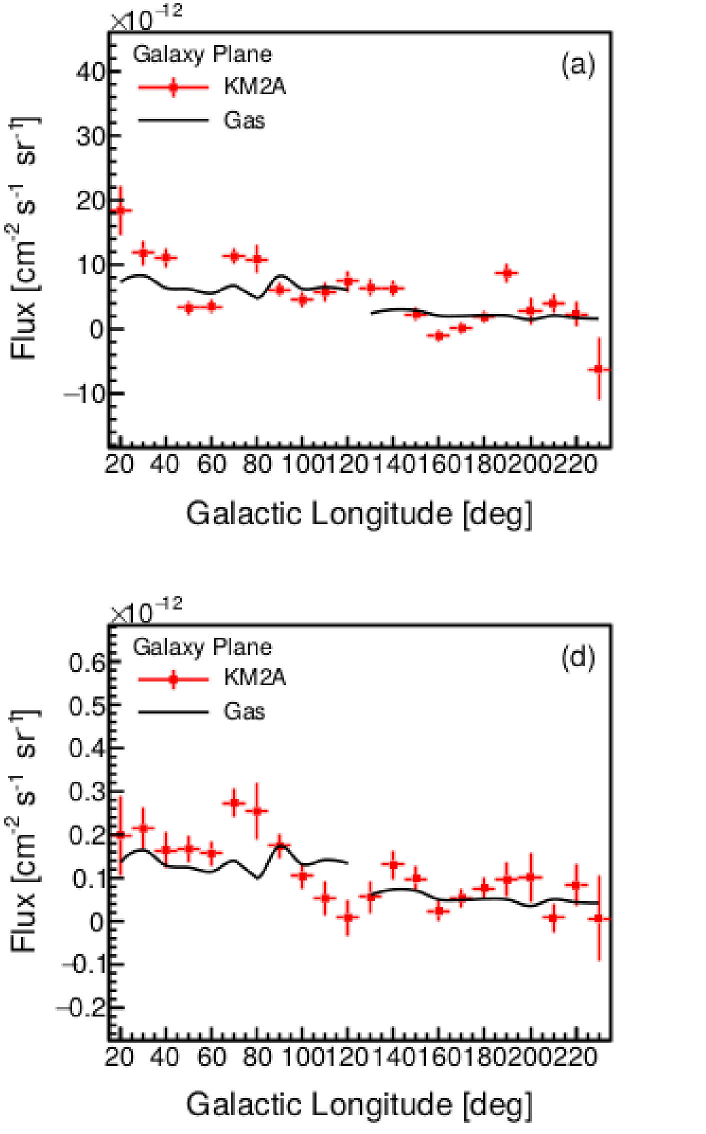


AS-γ Collaboration,
arXiv:2104.05181

The sky at ~ 10 TeV – 1 PeV: Diffuse emission from LHAASO

LHAASO Collaboration, arXiv:2305.05372





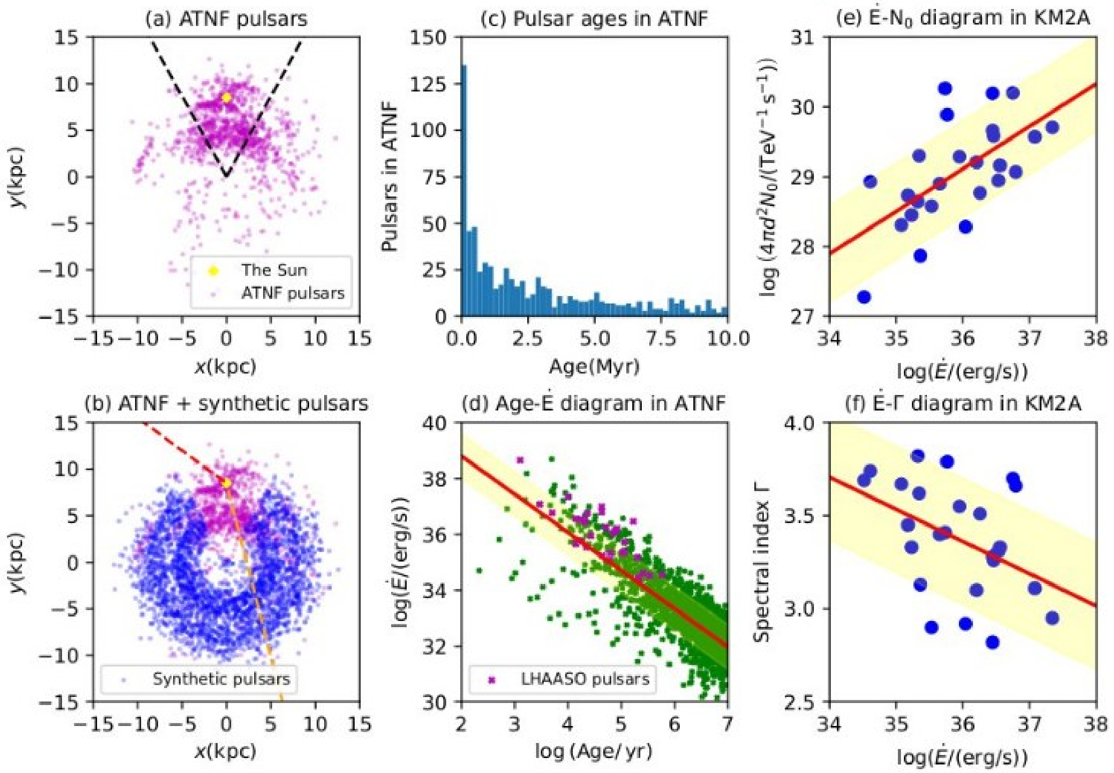
Emission in Galactic longitude does not follow target gas…

→ Leptonic origin?

→ Stochasticity of CR injection?

Impact of unresolved sources (PWNe)

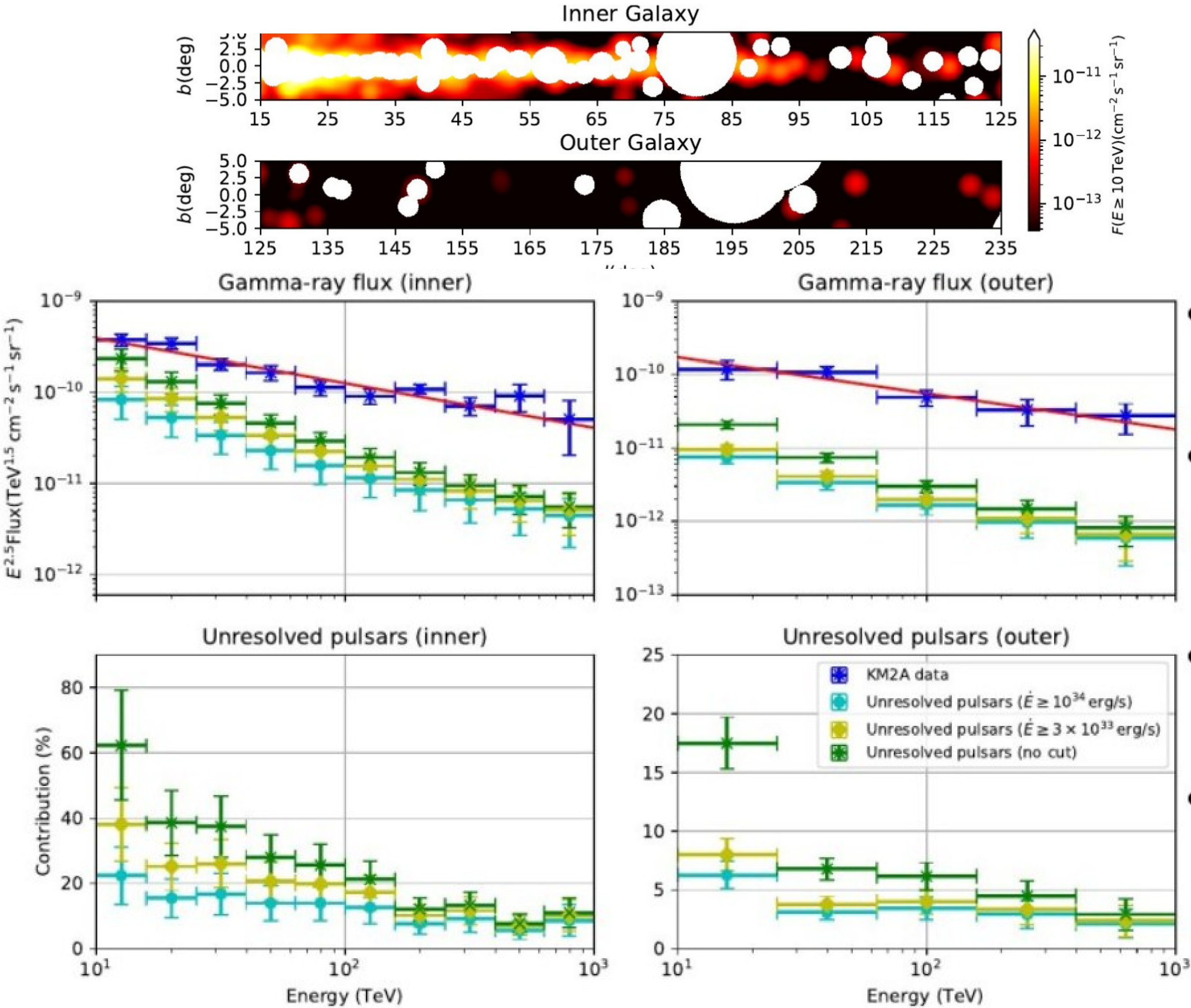
Kaci, Giacinti, Semikoz, ApJ Lett. 975, L6 (2024)



- Use ATNF catalog and complete it.
- Generate a VHE gamma-ray emission similar to that measured by KM2A for each source.
- Constrain the gamma-ray emission to be below KM2A sensitivity.
- Use the same masks as LHAASO.
- Compare the contribution of unresolved sources to the total flux measured by KM2A.

Upper Limits unresolved PWNe/halos

Kaci, Giacinti, Semikoz,
ApJ Lett. 975, L6 (2024)



Diffuse VHE γ-ray emission from discrete sources

SIMULATION:

Kaci & Giacinti, JCAP 01, 049 (2025)

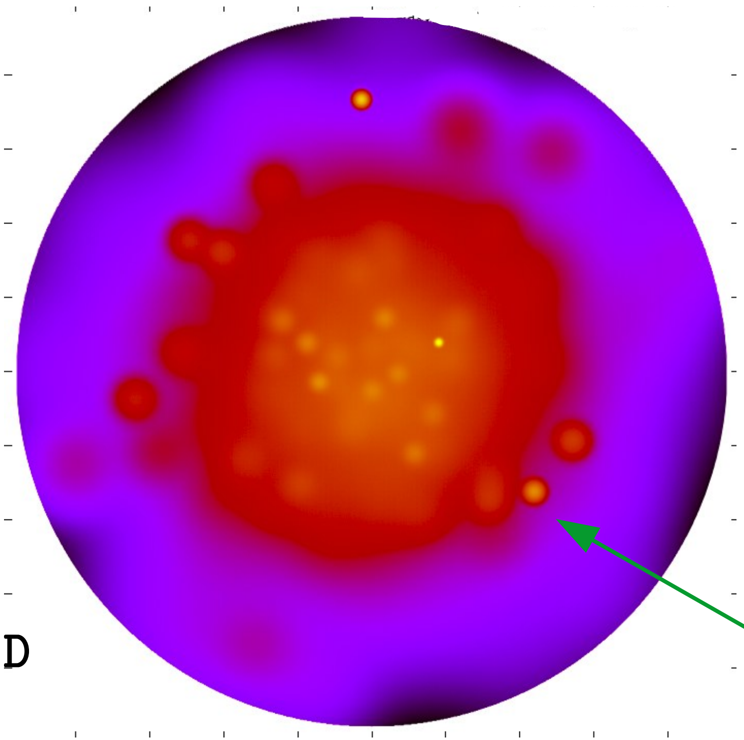
Isotropic and homogeneous diffusion

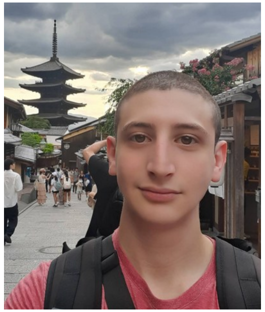
1) GALPROP-like (d=1/3) :

$$D(E) = 10^{28} D_{28} \left(\frac{R}{3GV}\right)^{\delta} cm^2/s$$

$$D_{28} = 1.33 \times \frac{H}{kpc}$$

2) Time-dependent (mimics self-confinement): 1/100 x D around sources for 10 kyr.





Samy Kaci

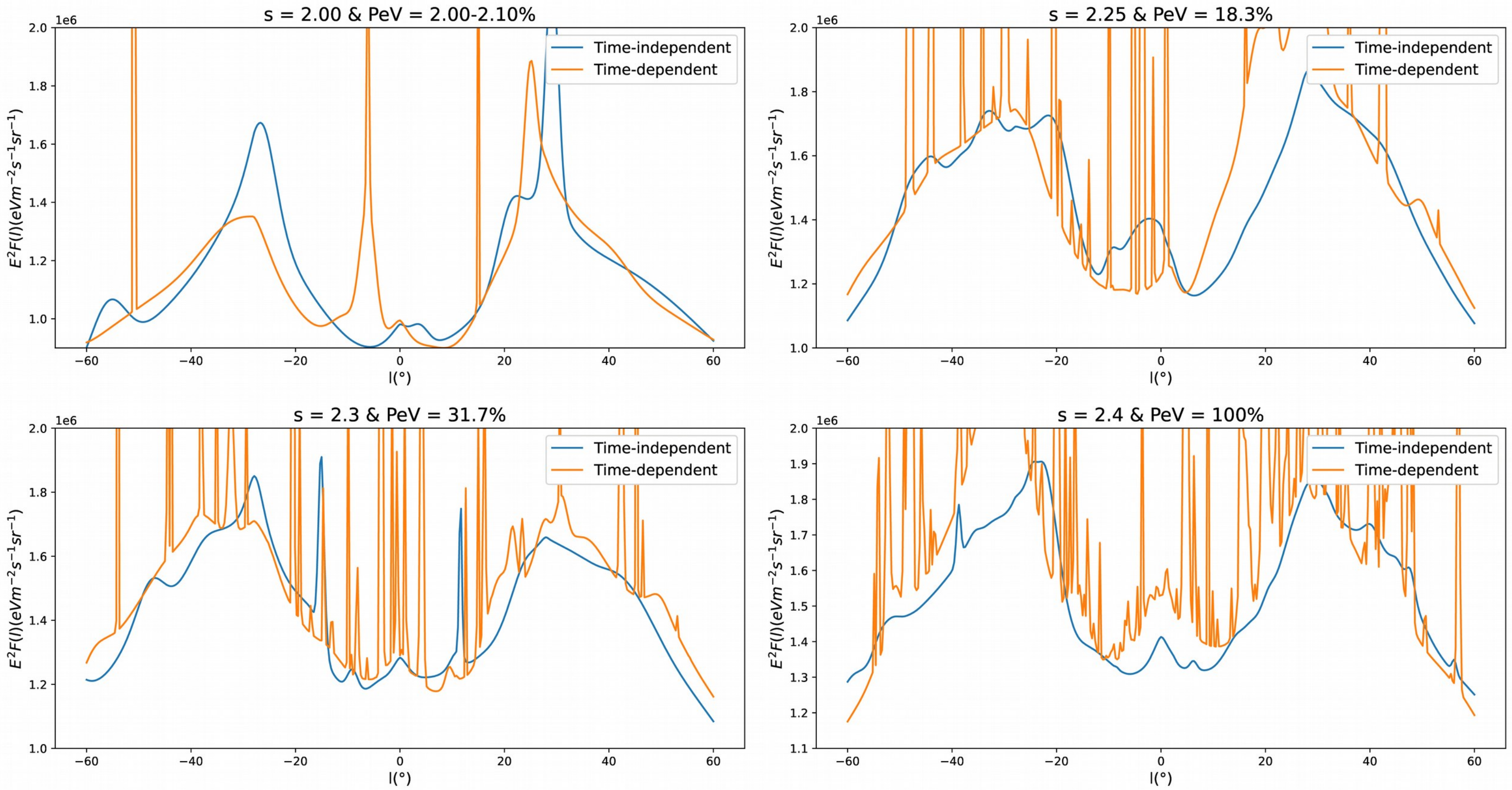
Cosmic-ray flux at Earth and B/C ratio satisfied

Discrete injection of cosmic rays

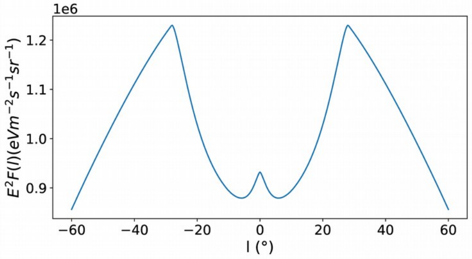
Slide from Samy Kaci

Clumps in the gamma-ray flux

Kaci & Giacinti, JCAP 01, 049 (2025)



Vs shape from Lipari & Vernetto(2018):



Diffuse gamma-ray flux clumpy at VHE

Microquasars as Galactic super-PeVatrons

Kaci, Giacinti et al., To be submitted (2025)

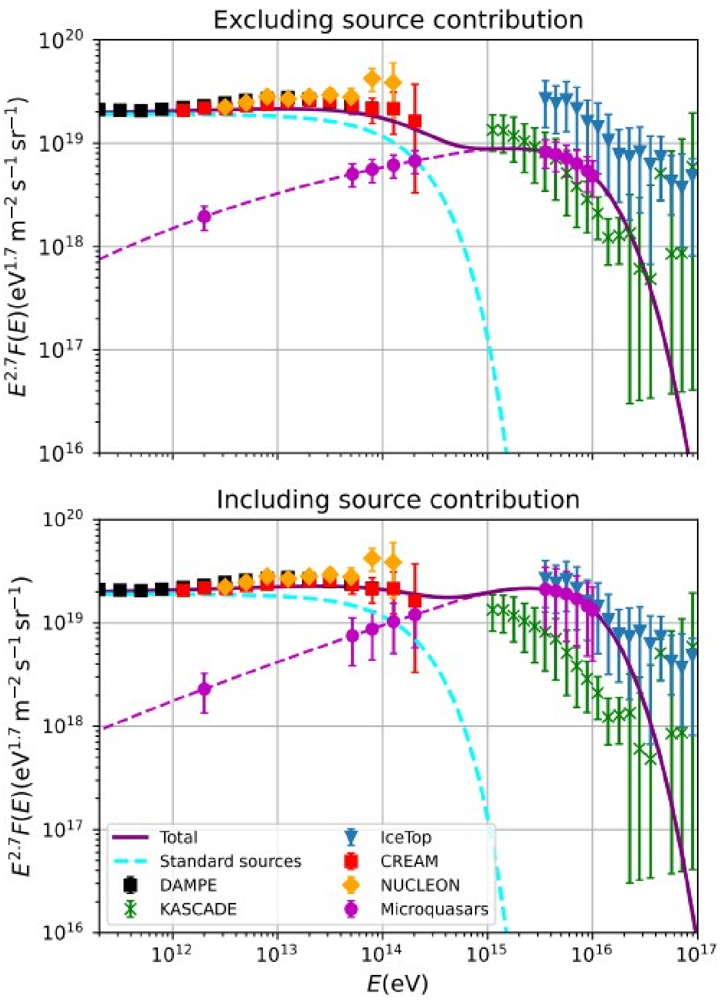
Microquasars: inj. -2 spectrum, cutoff at 10 PeV, 100 kyr lifetime, 0.1/kyr rate, power $10^{39\text{-}40}$ erg/s.

(Standard/SNR sources: -2.7 Power-law & exponential cutoff at 150TeV.)

-Up: Lower limit by excluding the existence of any active source within 4 kpc in the last 500 kyr.

- Down: No source younger than 100 kyr within 2.2 kpc. Only objects reported by BlackCat at r < 2.2 kpc, in the FoV of LHAASO and no detection by LHAASO.

→ Fits the knee and the 10TeV bump.

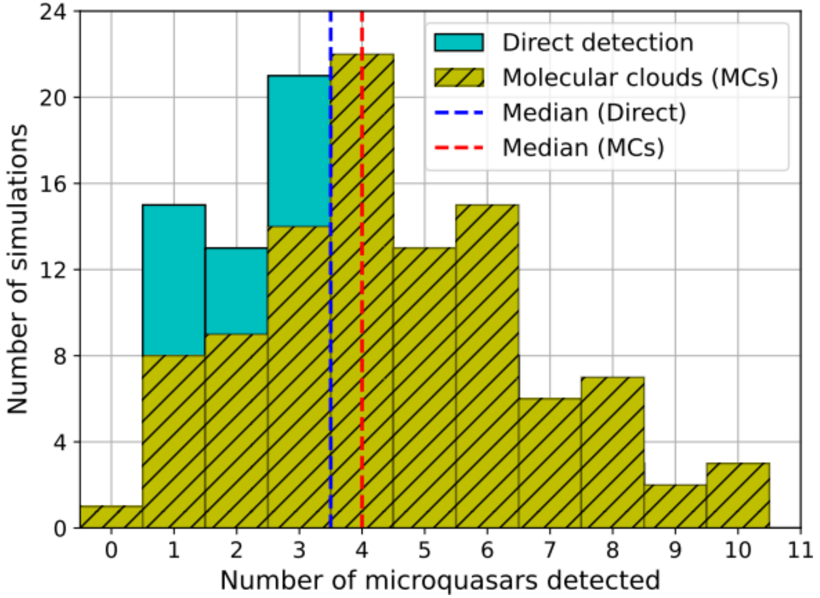


Microquasars as Galactic super-PeVatrons

Kaci, Giacinti et al., To be submitted (2025)

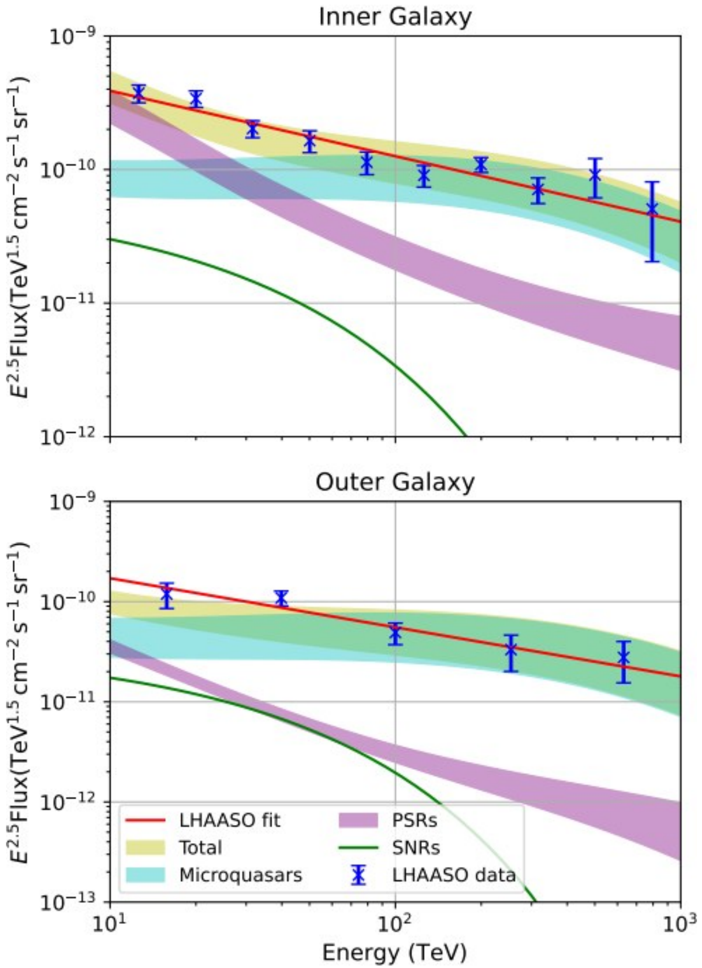
→ **POSTER 8**

On average 3-4 sources detectable:



molecular clouds similar to that around SS433
(uniform sphere with r = 20pc and n=30cm^{-3})

Contribution to the diffuse emission:



Conclusions & Perspectives

→ **Particle acceleration possible in equatorial region of pulsar wind termination shocks**

→ **TeV halos** constrain CR propagation around middle-aged pulsars

→ **Best fit** parameters for **Geminga**

→ **“Mirage” sources, asymmetric γ-ray sources**: **Can explain “weird” LHAASO sources**

→ **Clumpy diffuse γ-ray background at VHE**

→ **Unresolved PWNe/halos**: Minor contribution to LHAASO diffuse emission at > a few 10 TeV

→ ~ 10 powerful **microquasars can fit CR spectrum and LHAASO diffuse γ-ray emission.**